

NOVA LIGHT CURVES FROM THE SOLAR MASS EJECTION IMAGER (SMEI). II. THE EXTENDED CATALOG

R. HOUNSELL^{1,2}, M. J. DARNLEY³, M. F. BODE³, D. J. HARMAN³, F. SURINA^{3,4}, S. STARRFIELD⁵, D. L. HOLDSWORTH^{6,7}, D. BEWSHER⁷, P. P. HICK⁸, B. V. JACKSON⁹, A. BUFFINGTON⁹, J. M. CLOVER⁹, AND A. W. SHAFTER¹⁰ ² Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA

²Department of Astronomy, University of Illinois at Urbana–Champaign, 1002 W. Green St., Urbana, IL 61801, USA ³Astrophysics Research Institute, Liverpool John Moores University, IC2 Liverpool Science Park, Liverpool, L3 5RF, UK ⁴Department of Science and Technology, Chiangrai Rajabhat University, 80 Moo 9, Bandu, Muang, Chiangrai, 57100, Thailand

School of Earth and Space Exploration, Arizona State University, P.O. Box 871404, Tempe, AZ 85287-1404, USA

Astrophysics Group, Keele University, Staffordshire ST5 5BG, UK

⁷ Jeremiah Horrocks Institute, University of Central Lancashire, Preston PR1 2HE, UK

⁸ CAIDA UCSD/SDSC 9500 Gilman Dr., Mail Stop 0505, La Jolla, CA 92093-0505, USA

⁹ Center for Astrophysics and Space Science, University of California, San Diego, 9500 Gilman Drive 0424, La Jolla, CA 92093-0424, USA

Department of Astronomy, San Diego State University, San Diego, CA 92182, USA Received 2014 November 25; accepted 2015 November 30; published 2016 March 28

ABSTRACT

We present the results from observing nine Galactic novae in eruption with the Solar Mass Ejection Imager (SMEI) between 2004 and 2009. While many of these novae reached peak magnitudes that were either at or approaching the detection limits of SMEI, we were still able to produce light curves that in many cases contained more data at and around the initial rise, peak, and decline than those found in other variable star catalogs. For each nova, we obtained a peak time, maximum magnitude, and for several an estimate of the decline time (t_2) . Interestingly, although of lower quality than those found in Hounsell et al., two of the light curves may indicate the presence of a pre-maximum halt. In addition, the high cadence of the SMEI instrument has allowed the detection of low-amplitude variations in at least one of the nova light curves.

Key words: novae, cataclysmic variables – space vehicles: instruments – techniques: photometric

1. INTRODUCTION

Classical novae (CNe) belong to the cataclysmic variable (CV) class of objects. These are interacting close binary systems in which mass is transferred from a donor star to the surface of an accreting compact companion, leading to a variety of behaviors, the most noticeable of which is an eruption (see Bode & Evans 2008; Bode 2010, for reviews).

A typical CN system consists of a white dwarf (WD) primary and a cooler, lower-mass main-sequence secondary star (spectral type ranging from F to M; Darnley et al. 2012). As the secondary evolves, it fills its Roche lobe, allowing hydrogen-rich material to transfer at a rate of 10^{-11} – $10^{-9} M_{\odot}$ yr^{-1} via the inner Lagrangian point (L1) toward the WD and form an accretion disk. Material is deposited from the disk onto the surface of the WD, causing the bottom of the accreted layer on the WD to become electron degenerate. Within this degenerate layer hydrogen burning ensues, leading to a thermonuclear runaway (TNR) and ultimately the CN eruption. The total radiant output of a single CN eruption is in the range of 10⁴⁵-10⁴⁶ erg, with TN-outburst amplitudes of approximately 10-20 mag and absolute V-band magnitudes of $M_V = -10.7$ at maximum for the fastest and most luminous CNe (Shafter et al. 2009 and references therein). The resulting energy released is sufficient to expel the accreted envelope and drive mass loss $(10^{-5}-10^{-4} M_{\odot};$ Goss et al. 2011) at velocities of a few hundred to several thousand kilometers per second⁻¹. All CNe are thought to have repeat eruptions (although for some this may take up to $10^4/10^5$ yr; Hernanz 2005); however, a CN observed in TN outburst more than once is reclassified as a recurrent nova (RN); observed recurrence times range from 1 to 100 yrs (Darnley et al. 2014, 2015), or even as short as 6 months (Henze et al. 2015), with TN-outburst amplitudes

typically smaller than CNe (\approx 5–6 mag if the secondary is giant, 10-15 if a main-sequence star). The basic triggering and explosion mechanism of an RN is the same as for a CN, but there are some distinct differences in the physical properties of this subgroup. In order to reconcile the short quiescence period of an RN with a TNR, the WDs within these systems are believed to be hotter, more massive (close to the Chandrasekhar 1931, limit), and have higher mass accretion rates $(\sim 10^{-8} - 10^{-7} M_{\odot} \text{ yr}^{-1})$ than CNe (see Yaron et al. 2005; Wolf et al. 2013). The majority of RN systems also harbor evolved secondary stars rather than the typical main-sequence star of a CN system (see, e.g., Darnley et al. 2012).

As with most transients, CNe are traditionally classified according to their photometric and spectroscopic properties. Each nova has its own unique optical light curve. However, they do share a common idealized nova light curve (McLaughlin 1960). Most novae tend to rise rapidly to peak within 1-3 days. Owing to the transient nature of a nova, this initial rise has rarely been observed well enough to establish any classification regime, but the subsequent decline from maximum has. Novae tend to be classified according to the number of days that they take to decline n magnitudes from maximum; thus, they are divided into "speed classes." These classes were first introduced by Gaposchkin (1957), and often declines of two or three magnitudes are quoted (t_2 and t_3 , respectively).

During the initial rise, novae are thought to experience a premaximum halt (PMH) approximately one to two magnitudes below maximum (McLaughlin 1960). The duration of this halt appears to be related to the speed class of the nova, lasting a few hours for fast novae and a few days for slow novae. Until recently the PMH had only been observed for a few slow novae such as DQ Her, V450 Cyg (Gaposchkin 1957), V723 Cas, V463 Sct (Hachisu & Kato 2004, who present a new interpretation of the observed long PMHs), and V5558 Sgr (Tanaka et al. 2011), with no strong evidence to suggest that it was also present for the faster speed classes. Hounsell et al. (2010b, hereafter Paper I), however, found from Solar Mass Ejection Imager (SMEI) observations that for all very fast/fast novae (classical and recurrent) studied in detail in that paper, a PMH was detected with a duration consistent with their speed class. The PMHs of the very fast novae observed in Paper I consisted of three to four SMEI data points in which a temporary reversal of the light curve was suggested and its duration was calculated as the time between the first and third change in gradient of the slope. The PMH of the one fast nova detected (V1280 Sco) was represented by a plateau in the light curve and made up of seven SMEI data points. Here the PMH duration was taken as the time between the first and second change of gradient. If a slower nova were to have been detected within the SMEI data, we might expect a similar plateau.

Although there is now evidence for the existence of this phase of evolution in the nova light curve, no physical explanation for the PMH is currently accepted, although several theories exist. Williams (2012) suggests that a PMH lasting a day or more may be caused by an enhancement of mass loss from the secondary star, which then dominates over the initial WD ejecta. This changeover in the dominance of the two mass-loss sources could be the cause of the halt. Other work by Hillman et al. (2014) has used nova evolution simulations (via a hydrodynamic Lagrangian code) to create detailed light curves, and find that halts of long or short duration occur naturally within the burst. Examining the evolution of nova effective temperature (T_{eff}) , luminosity, and radius, they find that at a point just before expansion and mass loss occur, the $T_{\rm eff}$ deceases and the rising luminosity halts. Their work also found that this halt would in many cases be accompanied by a dip in the total luminosity. Such a dip is due to a temporary drop in the energy flux as convection in the expanding and thinning envelope ceases to be efficient near the envelope surface. Because of the decreasing opacity of the envelope, radiation soon dominates over convection, thus reversing the dip. Work by both Hillman et al. (2014) and Williams (2012) is WD mass dependent.

The majority of Galactic novae are still discovered by the amateur community. Over the past century various surveys have attempted to measure the Galactic rate; however, results are greatly affected by temporal and spatial coverage, selection effects, and interstellar extinction. Based on an extrapolation of the observed nova rates in the solar neighborhood (assumed complete to m < 2), Shafter (1997) estimated a Galactic nova rate of approximately 35 yr^{-1} ; this is supported by Darnley et al. (2006), with a value of 34^{+15}_{-12} yr⁻¹. Of these, an average of roughly one CN per year has been *observed* to reach $m_V = 8$ or brighter (see Figure 2 of Shafter 2002). It is well known that novae with the highest peak bolometric luminosity fade the most rapidly and are thus often missed; historical observations are therefore clearly incomplete at $m_V = 8$, and the actual number of novae reaching this brightness is expected to be significantly higher ($\sim 5 \text{ yr}^{-1}$; see Shafter 2002; Warner et al. 2008, and discussion in Paper I).

With the advent of all-sky imaging facilities, both ground and space based, there is new hope for detecting a more complete sample of Galactic novae. Although the detection of transient events may not be the original science objective of these missions, their archives hold a wealth of data on many events. These observations may contain great detail and provide data on many previously poorly examined and understood phases of evolution. Examination of these archives is therefore exceptionally important. One such space-based allsky mission is the SMEI; its usefulness in observing novae was well documented in Paper I, Hounsell et al. (2010a, 2011, 2012), Surina et al. (2014), and Darnley et al. (2013).

SMEI is a high-precision white-light differential photometer (Buffington et al. 2006, 2007) based on board the *Coriolis solar satellite*. SMEI was in operation from 2003 January (Eyles et al. 2003; Jackson et al. 2004) until 2011 September. The instrument consists of three baffled CCD cameras each with a $60^{\circ} \times 3^{\circ}$ field of view, combining to sweep out nearly the entire sky with each 102-minute orbit of the spacecraft (Hick et al. 2007). The peak throughput of the instrument is at approximately 700 nm with an FWHM ~ 300 nm. SMEI is able to reliably detect brightness changes in point sources down to at least 8th magnitude.

SMEI was originally designed to map out large-scale variations in heliospheric electron densities by observing the Thomson-scattered sunlight from solar wind electrons (Jackson et al. 2004). In order to isolate the faint Thomson-scattered sunlight, the much larger white-light contributions from the zodiacal dust cloud, the sidereal background, and individual point sources (bright stars and planets) were determined and removed (see Hick et al. 2007, for further details). Thus, brightness determination of point sources is a routine step in the SMEI data analysis and as such has led to the production of detailed light curves for many bright objects including variables (see, e.g., Buffington et al. 2006; Spreckley & Stevens 2007; Tarrant et al. 2007, 2008a, 2008b, 2008c; Clover et al. 2011; Goss et al. 2011; Surina et al. 2014; Paper I).

This paper presents results from the examination of nine additional Galactic novae observed by SMEI. These novae are fainter than the four presented in Paper I and indicate the limit of SMEI's ability to detect brightness changes at fainter magnitudes and in crowded regions. In Section 2 we explain how data were obtained and analyzed. Section 3 presents our results, and a discussion are conclusions are provided in Section 4.

2. DATA ANALYSIS

The SMEI database contains a catalog with the names, coordinates, and discovery magnitudes of 62 Galactic CNe and 3 RNe, with eruptions dating between 2003 and 2011. The photometry of each nova in this paper was obtained using this catalog and an extended iterative least-squares fit of the pointspread function (PSF) as described in Hick et al. (2005, 2007) and Section 2 of Paper I. Zodiacal and sidereal background light contributions were also considered during the fitting stage. The remaining 51 novae not examined as part of this paper, Surina et al. (2014), or Paper I either were too faint for detection or resided in fields far too crowded for reliable results to be obtained.

As noted in Paper I, the SMEI PSF has a full width of approximately 1° and is highly asymmetric with a "fish-like" appearance. Owing to this large size, an object of interest is considered crowded when it lies less than one PSF width from another bright object (typically 6th magnitude or brighter);

 Table 1

 Derived Light-curve Parameters of 14 Novae Using Data from the Solar Mass Ejection Imager

Name	Time of Maximum (yyyy/mm/dd)	Peak SMEI Magnitude	t ₂ (days)	Speed Class ^a	Pre-max Duration (days) ^b	Pre-max Mean Magnitude	$\Delta m_{\rm SMEI}$ From Halt to Peak ^c	Δt From Halt to Peak (days)	Wing Radius (deg)
V1187 Sco	$2004/08/3.77^{+0.07}_{-0.04}$	6.95 ± 0.04	$10.10^{+0.43d}_{-0.39}$	Very fast					1.3
V2467 Cyg	$2007/03/16.56\pm 0.04$	6.30 ± 0.03	$5.68^{+0.61d}_{-0.53}$	Very fast					1.3
V458 Vul	$2007/08/13.66\pm0.04$	7.94 ± 0.07		Fast					auto-wing
V597 Pup	$2007/11/14.68^{+0.04}_{-0.35}$	6.91 ± 0.04	$2.81^{+1.0}_{-0.90}$ ^d	Very fast	0.21	8.07 ± 0.14	1.16	0.96	1.3
V459 Vul	$2007/12/28.11\substack{+0.11\\-0.07}$	6.59 ± 0.04	~19.4 ^e	Fast					1.2
V2491 Cyg	$2008/04/10.89\pm0.04$	7.36 ± 0.05		Very fast					1.2
QY Mus	$2008/09/28.63 \pm 0.2$	6.93 ± 0.04		Moderately fast					1.2
V5580 Sgr	2008/11/30.85	7.01 ± 0.04		••••					auto-wing
V5583 Sgr	$2009/08/7.08\pm0.04$	6.94 ± 0.05		Very fast	0.21	7.64 ± 0.03	0.71	1.10	1.4
T Pyx ^f	$2011/05/12.22\pm0.04$	6.33 ± 0.03		Slow	~10	~8.1	~1.8	14.6	
RS Oph	$2006/02/12.94 \pm 0.04$	3.87 ± 0.01	7.9	Very fast	0.14	4.50 ± 0.05	0.63	0.49	1.25
V1280 Sco	$2007/02/16.15\pm0.04$	4.00 ± 0.01	21.3	Fast	0.42	5.231 ± 0.003	1.23	0.49	1.25
V598 Pup	$2007/06/6.29 \pm 0.04$	3.46 ± 0.01	4.3	Very fast	0.28	5.2 ± 0.1	1.74	2.19	1.4
KT Eri	$2009/11/14.67\pm0.04$	5.42 ± 0.02	6.6	Very fast	0.14	6.04 ± 0.07	0.63	0.71	1.25

Notes. The last five novae listed represent objects reported within Hounsell et al. (2012), Surina et al. (2014), and Paper I. Novae in this paper tend to be fainter than those in Paper I, and as such we were unable to obtain in most cases all parameters listed in the table.

^a V1187 Sco—Lynch et al. (2006), V2467 Cyg—Lynch et al. (2009), V458 Vul—Wesson et al. (2008), V597 Pup—Naik et al. (2009), V459 Vul—Poggiani (2010), V2491 Cyg—Munari et al. (2011) and Darnley et al. (2011), QY Mus—Schwarz et al. (2011), V5583 Sgr—Schwarz et al. (2011).

^b Here the duration of the halt is taken to be the time between the first and third change in gradient of the rising light curve for very fast novae. For fast and moderately fast it is the first and second change.

 $^{c}\Delta m_{\rm SMEI}$ from halt to peak is calculated using the mean magnitude of the PMH.

^d Based on a linear fit of the initial decline data from the SMEI nova light curve (see text for further details).

^e Using linear extrapolation between the first and last data point of the SMEI light curve (see text for further details).

^f Here all data have been taken from Surina et al. (2014). Within this work the PMH was suggested to have lasted for ≈ 10 days with magnitudes varying between 7.7 and 8.5; hence, a mean magnitude of 8.1 was quoted in the above table. The time between halt and peak was also taken, in this case, from the end of the PMH phase.

such crowding within the SMEI data is commonplace. To combat this issue, simultaneous fitting of multiple objects can be initiated where possible (stellar separation $>0^{\circ}.75$) and contamination of the object of interest reduced.

In order to achieve the best fit and produce the most reliable nova light curve, the surrounding region of each object required assessment for levels of potential contamination due to crowding, cosmic-ray hits, and finally their proximity to the sunward and antisunward masks (see Buffington et al. 2006, 2007, for a description of these masks). If the nova was in a crowded field, a simultaneous fit was conducted and the area used to sample the surrounding stellar region (wing radius) from the PSF centriod reduced. The reduction of the wing radius also reduces errors if the object is located near a masked region. For these reasons multiple light curves were generated for each nova using an *auto_wing*¹¹ radius and wing radii of 1°.2, 1°.3, and 1°.4. The resulting photometry files were then assessed on several criteria: (1) correlation of the fitted PSF to the model, via the sample Pearson correlation coefficient (r); (2) the number of pixels used to define the PSF fitting area, which is often an indictor of masks or image defects (npsf); (3) variation of the background fitted value; (4) deviation of the R. A. and decl. from the catalog position, where jumps can be an indication of the fitting of a residual from the poor subtraction

of a neighboring star; this becomes more of an issue as the nova fades.

Within Paper I, points possessing $r \ge 0.5$ were deemed reliable. However, as the novae presented here are much fainter, a less stringent cutoff of 0.4 was found to be acceptable. A 1 σ threshold was applied to the other selection criteria listed above (in some cases a 3σ cutoff was applied if the object was detected only a few times by the instrument, or if the initial R. A. and decl. were uncertain). Using the filtered file that possessed the highest number of "valid" points, the flux of the nova was then converted into an unfiltered SMEI apparent magnitude (m_{SMEI}) and its error contribution calculated from photon counting statistics.

3. RESULTS

As noted above, the novae presented in this paper are much fainter than those given in Paper I, all of which possessed a peak $m_{\text{SMEI}} < 5.5$. The objects within this paper have peak magnitudes between 6 and 8; fainter than this detection is unreliable. The light curves derived, although more noisy than those in Paper I, have nonetheless provided precise measurements of peak time, eruption magnitude, and in many cases a value for t_2 . Table 1 summarizes our main findings, and the individual novae are discussed below. Novae that possess the most interesting data are presented first, with the remaining given in order of eruption date. Where possible a light curve for each nova has been presented, with a legend for all data given in Figure 1. An Appendix tabulating the available photometry for each nova examined is included at the end of the paper.

¹¹ This is defined by the following equation: $auto_wing = 0.1 \times (3 - m_{SMEI}) + 1.4$ (degrees), where m_{SMEI} is the magnitude listed for the nova within the catalog and is often the discovery magnitude only. For novae of interest $auto_wing$ is approximately 1°.

THE ASTROPHYSICAL JOURNAL, 820:104 (27pp), 2016 April 1

3.1. V597 Puppis

Nova V597 Pup ($\alpha = 08^{h}16^{m}18^{s}.01$, $\delta = -34^{\circ}15'24''_{v}1$; J2000) was discovered in TN outburst by Pereira et al. (2007) at a visual magnitude of 7.0 on 2007 November 14.23 UT (MJD 54,418.23), reaching a peak visual magnitude of $m_{V} = 6.4$ on 2007 November 14.48 UT (MJD 54,418.48, green cross within Figure 2). The nova then declined rapidly with a $t_{2} = 2.5$ days (Naik et al. 2009), making it one of the fastest novae recorded. Naik et al. (2009) classified the nova as



Figure 1. Legend for all data in following figures.

an He/N-type with a WD close to the Chandrasekhar limit. A pre-eruption detection is found within the Digitized Sky Survey (Pereira et al. 2007) with a source at $V \sim 20$, which coincides with the nova position. Continued monitoring of the object by Warner & Woudt (2009) revealed the nova as an intermediate polar (IP) in the orbital period gap ($P_{orb} = 2.67$ hr), with a rotational period of 8.7 minutes. Observation of the object a year after eruption also revealed the presence of a deep secondary eclipse caused by the passage of the optically thick accretion disk in front of the irradiated side of the secondary star. The object is considered unique as it is the first CV found to have this deep an eclipse.

The light curve created using the SMEI data set covers the latter part of the initial rise, peak, and early decline (see Figure 2). There does appear to be one data point on 2007 November 7.04 UT (MJD 54,407.04) at $m_{\text{SMEI}} = 9.57 \pm 0.13$; however, the r of this point is only 0.34 and thus not reliable. All data after this point and up to the initial rise possess an r less than 0.3 and are therefore omitted from Figure 2 and further discussion. Examination of the initial rise hints at the possibility of a PMH starting on 2007 November 13.63 UT (MJD 54,417.63) and lasting several hours with a mean $m_{\rm SMEI} = 8.07 \pm 0.14$ (defined as the mean magnitude over the duration of the halt; quoted error is the rms scatter). As in Paper I, the duration of the halt for this object is taken to be the time between the first and third changes in gradient of the rising light curve and is appropriate for the speed class of the nova. It should be noted, however, that the first two points of this halt have r < 0.4 (actual values are 0.39 and 0.35, respectively; this lower r is due in part to the limiting magnitude of SMEI), and



Figure 2. Optical light curves of V597 Pup. SMEI data with a lower *r* than optimal have been included for completeness. Nova Atlas data that coincide with the SMEI light curve have been included along with data from the AAVSO (see Figure 1 for data legend). Green and red crosses represent the peak *V*- and *R*-band data given within Pereira et al. (2007). The inset shows the final rise, peak, and initial decline of the nova in the SMEI data only, which may show evidence for a PMH around MJD 54,417.63.

as such the reality of this PMH can not be confirmed. There is then a gap in the data lasting several hours due to the SMEI weekly calibration, after which the nova is seen to reach its maximum intensity of $m_{\text{SMEI}} = 6.91 \pm 0.04$ on 2007 November 14.68^{+0.04}_{-0.35} UT (MJD 54,418.68). This coincides better with the R-band magnitude of 6.7 on 2007 November 14.48 UT (MJD 54,418.48) given in Pereira et al. (2007; red cross within Figure 2) than the V-band data quoted above. Note, however, that it is possible that the peak nova brightness actually occurred in the data gap caused by the weekly calibration. This uncertainty has been included in the quoted time of maximum. The decline of the nova is shown by SMEI to be exceptionally fast, but given the scatter off the data around 8/9 magnitude, a direct measurement of t_2 could not be made. However, assuming that the data can be modeled by a simple linear decline (as is often seen in the early stages of nova light curves), we derived an approximate t_2 time of $2.81^{+1.0}_{-0.9}$ days (error quoted is statistical only and was obtained using the uncertainties of the fitted gradient and intercept). This is similar to the value given by Naik et al. (2009), confirming the very fast classification. Within Figure 2, the SMEI light curve is also compared to AAVSO data (V band) and data (B, V, R, and I band) from the STONY BROOK/SMARTS Atlas of (mostly) Southern Novae¹² (hereafter referred to as the Nova Atlas; see Walter et al. 2012, for a full description) within the same time frame.

3.2. V5583 Sagittarii

Nova V5583 Sgr ($\alpha = 18^{h}07^{m}07^{s}.67$, $\delta = -33^{\circ}46'33''.9$; J2000) was observed by both SMEI and the *Solar TErrestrial Relations Observatory (STEREO)* Heliospheric Imager (Eyles et al. 2009). A review of *STEREO* observations for this object can be found in Holdsworth et al. (2014).

V5583 Sgr is a very fast CN that was discovered by Nishiyama et al. (2009) at a magnitude of 7.7 on 2009 August 6.49 UT (MJD 55,049.49); it then rose to maximum on 2009 August 7.57 UT (MJD 55,050.57) at $m_V = 7.43$. An approximate t_2 time of 5 days is given by Schwarz et al. (2011).

The nova was detected by SMEI on 2009 August 1.16 UT (MJD 55,044.16) at $m_{\rm SMEI} = 8.60 \pm 0.10$. There is then a gap in the data due to a combination of low *r* and *npsf* values lasting until August 5.67 UT (MJD 5,548.67), after which it then rises (Figure 3) to a peak magnitude of $m_{\rm SMEI} = 6.94 \pm 0.05$ on 2009 August 7.08 ± 0.04 UT (MJD 55,050.08). The timing of this peak coincided nicely with that presented in Holdsworth et al. (2014; the *STEREO* peak is given as $m_{\rm H\,I} = 7.18 \pm 0.06$ on 2009 August 7.04 ± 0.17 UT, MJD 55,050.04); however, the timing of both the SMEI and *STEREO* peaks is earlier than that given by Nishiyama et al. (2009), suggesting that the maximum was previously missed.

A PMH may be present in the SMEI data around 2009 August 5.98 UT (MJD 55,048.98) at an average $m_{\text{SMEI}} = 7.64 \pm 0.03$ (defined as the mean magnitude over the duration of the halt; quoted error is the rms scatter). Unfortunately, the feature is derived from data with a much lower *r* than optimal (*r* values are 0.23, 0.26, and 0.26, respectively) and is therefore unreliable. However, it should be noted that Holdsworth et al. (2014) also find evidence of PMHs around this time. Holdsworth et al. (2014) define a PMH as a decrease in the brightness during the rise phase of the nova eruption and detect three possible events, the first on MJD 55,048.88 \pm 0.08 at $m_{\rm H\,I}$ = 8.38 \pm 0.09, the second on MJD 55,049.13 \pm 0.08 at $m_{\rm H\,I}$ = 8.05 \pm 0.10 (which, when taking into consideration timing definitions, coincides with the SMEI PMH), and the third on MJD 55,049.29 \pm 0.08 at $m_{\rm H\,I}$ = 7.89 \pm 0.08 (see Figure 4). This third halt is also seen in the SMEI light curve at an average $m_{\rm SMEI}$ = 7.18. However, it is only ~0.25 mag below maximum and is not within the range that has generally been proposed for a PMH.

Approximately 1 day after maximum, the object then seems to enter a "plateau" phase that lasts for 1.6 days with an average $m_{\text{SMEI}} = 7.56$, after which it declines. Again it should be noted that these data have a correlation coefficient much less than optimal ($0.1 \le r < 0.3$) and are less reliable. The discrepancy between the SMEI and *STEREO* light curves especially at late times is clearly evident and due to the reduced quality of the SMEI data (r < 0.3) throughout the nova eruption. Therefore, the r < 0.3 SMEI data are less reliable and complementary to the *STEREO* data when obtained simultaneously.

3.3. V1187 Scorpii

Nova V1187 Sco ($\alpha = 17^{h}29^{m}18^{s}.81, \delta = -31^{\circ}46'01''.5;$ J2000) was discovered prior to peak at $m_V = 9.9$ on 2004 August 2.07 UT (MJD 53,219.07) using data from the All Sky Automated Survey (ASAS)-3 patrol (Pojmanski et al. 2005). It subsequently rose to a maximum magnitude of $m_V = 7.42$ on 2004 August 3.58 UT (MJD 53,220.58; Yamaoka 2004; this is the peak of the V-band AAVSO light curve given in Figure 5). The initial decline gave $t_2 = 8.7$ and $t_3 = 15$ days (Lynch et al. 2006), classifying V1187 Sco as a very fast nova. Near-IR spectroscopic observations of the object by Lynch et al. (2006) indicated the development of a nova explosion on an O Ne WD, which did not form dust before entering its nebular phase. The emission lines found within the spectra had complex, double-peaked profiles. Using the H1 double emission lines, the nova ejecta were modeled as a ring or partial sphere-like emitting region. An extinction of $A_V = 4.68 \pm 0.24$ was derived also using O1 lines in combination with the optical spectra.

The SMEI data give the peak brightness of nova V1187 Sco as $m_{\text{SMEI}} = 6.95 \pm 0.04$ on 2004 August $3.77^{+0.07}_{-0.04}$ UT (MJD 53,220.77). This is 0.47 mag brighter and 0.19 days later than the *V*-band peak given in the lower time resolution results of Yamaoka (2004). Within Figure 5, the SMEI light curve is compared to *V*-band AAVSO data and *I*-band data from the Nova Atlas within the same time frame.

A substantial portion of the initial nova rise is missing from the SMEI light curve; however, some data before peak are presented. As the nova is of the very fast classification, one would expect the rise to maximum to also be fast; however, within the SMEI data we are finding the opposite effect. This apparent slow rise is most likely due to contamination of the data from several close bright neighboring stars, despite the conduction of simultaneous fitting and the fitting of additional bright stars in the larger surrounding region. The apparent plateau/dip before the final rise to peak may therefore not be a PMH; instead, it is probably residual light from a neighboring star. Data points between July 29 and 31 (MJD 53,215–53,217) were removed owing to large variations in the R.A. and decl. on fitting, suggesting a fit of residuals rather than the nova. However, the gap seen around the 2nd of August (MJD 53,219)

¹² http://www.astro.sunysb.edu/fwalter/SMARTS/NovaAtlas/



Figure 3. Optical light curves of V5583 Sgr. SMEI data with a lower r than optimal have been included for completeness. Nova Atlas data that coincide with the SMEI light curve have been included along with data from the AAVSO (see Figure 1 for data legend).



Figure 4. Optical light curves of V5583 Sgr. SMEI data with a lower r than optimal have been included for completeness. The *STEREO* H I data are given by orange points, the uncertainty of which is 0.06 mag above 10th magnitude. The dotted black lines indicate PMHs suggested in *STEREO* data (Holdsworth et al. 2014), the second of which coincides with that suggested by the SMEI light curve around MJD 55,048.98.

is due to an SMEI anomaly during which the instrument was forced to shut down and restart.¹³ The noisiness of the light curve especially during decline is also likely due to contamination from neighboring bright stars. However,

there does appear to be some structured oscillation that may be real.

The SMEI data shown in Figure 5 do not cover the magnitude range required for the determination of a t_2 time. However, on application of a liner fit to the initial decline of the nova using all data where *r* is greater than 0.3, a t_2 of $10.10^{+0.43}_{-0.39}$ days was derived (error quoted is statistical only and obtained using the uncertainties of the fitted gradient and intercept). This

 $[\]frac{13}{13}$ An SMEI anomaly means that the satellite underwent an impromptu manipulation, causing the system to go offline.



Figure 5. Optical light curves of nova V1187 Sco. Nova Atlas data that coincide with the SMEI light curve have been included along with data from the AAVSO (see Figure 1 for data legend). It is likely that the SMEI data presented here are contaminated by light from neighboring bright stars, and as such should be treated with caution.

is slightly larger than the result presented in Lynch et al. (2006), but agrees (within errors) with the very fast nova classification.

3.4. V2467 Cygni

Nova V2467 Cyg ($\alpha = 20^{h}28^{m}12^{s}47$, $\delta = +41^{\circ}48'36''_{*}4$; J2000) was discovered by A. Tago at $m_{V} = 7.4$ (Nakano et al. 2007c) on 2007 March 15.79 UT (MJD 54,174.79). It then rose to a peak magnitude of 6.7 on 2007 March 16.77 UT (MJD 54,175.77) and declined with a $t_{2} = 7.3$ (Lynch et al. 2009). An early spectrum of the nova obtained on 2007 March 16.8 UT (MJD 54,175.8) indicated an expansion velocity of ~1200 km s⁻¹ and an Fe II-type nova classification (Munari et al. 2007).

The transition phase of the object was seen to start in April of 2007 after fading approximately 4 mag. Within this phase, six quasi-periodic oscillations were observed with periods from 19 to 25 days and amplitudes of ≈ 0.7 mag. Swierczynski et al. (2010) proposed that the period found within the optical light curve and changes found in the subsequent X-ray detections could only be explained if the system were an IP.

The SMEI light curve for nova V2467 Cyg can be seen in Figure 6. The data show a gap between the first detectable point and the initial rise of the nova. Examination of data within this period suggests that the SMEI pipeline struggled to find a point source owing to a combination of noise from nearby objects and the overall faintness of the nova at this time. The initial point seen in Figure 6 is only just detectable above this background noise. The rise of the nova light curve starts on 2007 March 14.52 UT (MJD 54,173.52) and is very steep, rising 2.8 mag in just under 2 days. Within this rise to peak no PMH is found, possibly owing to a lack of data around $m_{\text{SMEI}} = 7.5$. The peak magnitude of the nova is given as $m_{\text{SMEI}} = 6.30 \pm 0.03$ on 2007 March 16.56 \pm 0.04 UT (MJD 54,175.56); this is 0.40 mag brighter and 0.21 days

earlier than the V-band peak given in Nakano et al. (2007c) and suggests that it was not covered within their data. Once again the large scatter of the SMEI data on decline prevents a direct measurement of t_2 . Assuming that any scatter in the data is due to noise and is not inherent to the intrinsic behavior of the nova, a linear fit of the initial decline yields an approximate t_2 time of $5.68^{+0.61}_{-0.53}$ days, classifying the nova as very fast. This is a sharper decline than that observed by Lynch et al. (2009). During this decay, there are several gaps in the SMEI data. The first takes place during March 16 (MJD 54,175) and is due to a masking issue that lasts for less than a day; the second occurs on March 18 (MJD 54,177) and is due to many factors, including engineering work, pre-annealing calibration, hot annealing, and an instrumental anomaly. This gap lasted for several days.

3.5. V458 Vulpeculae

Nova V458 Vul ($\alpha = 19^{h}54^{m}24$ °.61, $\delta = +20^{\circ}52'52''6$; J2000) was discovered in TN outburst by Nakano et al. (2007a) at an apparent magnitude of 9.5 on 2007 August 8.54 UT (MJD 54,320.54), reaching its peak visual magnitude soon after at $m_V = 8.1$. The nova's t_3 time is given as 21 days, and it is therefore classified as a fast nova (Wesson et al. 2008). The decline of the object was disrupted by two re-brightenings, the first occurring 20 days after maximum. Spectra taken after the eruption revealed the nova to be of the hybrid spectral class (Poggiani 2008), with an early Fe II-type spectra and evolving to He/N. V458 Vul is an interesting nova as it occurred within a planetary nebula, and as such light echoes of the burst within the surrounding material are seen (Wesson et al. 2008).

The SMEI light curve and corresponding AAVSO data of this nova are displayed in Figure 7. Unfortunately, SMEI covers only 8 days around the nova peak, as on other dates the nova is too faint for detection. What is most striking about this



Figure 6. Optical light curves of V2467 Cyg indicating the peak of the nova and its initial decline. SMEI data with a lower r than optimal have been included for completeness. AAVSO data that coincide with the SMEI light curve have been included (see Figure 1 for data legend). The two green crosses represent data from Nakano et al. (2007c). The inset shows the final rise, peak, and initial decline of the nova in the SMEI data only.



Figure 7. Optical light curves of V458 Vul indicating the multiple peaks of the nova. AAVSO data are displayed for comparison (see Figure 1 for data legend). There is a large color difference between blue and red filters.

light curve is that there are multiple flaring events during the initial decline (these may be akin to those seen in U Sco; Schaefer et al. 2011). There are two bright peaks separated by a fainter third; the magnitude and times of each peak are given as $m_{\text{SMEI}} = 8.04 \pm 0.07$ on 2007 August 9.93 \pm 0.04 UT (MJD 54,321.93), $m_{\text{SMEI}} = 8.30 \pm 0.08$ on 2007 August $11.48^{+0.08}_{-0.18}$

UT (MJD 54,323.48), and $m_{\text{SMEI}} = 7.94 \pm 0.07$ on 2007 August 13.66 \pm 0.04 UT (MJD 54,325.66). In comparison, the peak of the *R*-band light curve occurs on 2007 August 13.31 UT (MJD 54,325.31) at $m_R = 7.87 \pm 0.01$, with the peak of the *I* on 2007 August 13.87 UT (MJD 54,325.87) at $m_I = 7.35 \pm 0.01$.



Figure 8. Optical light curves of V459 Vul. The SMEI light curve indicates oscillations during the decline from maximum. Data from the AAVSO are displayed for comparison (see Figure 1 for data legend).

3.6. V459 Vulpeculae

Nova V459 Vul ($\alpha = 19^{h}48^{m}08^{s}.87, \delta = +21^{\circ}15'26''.8;$ J2000) was discovered independently by Hiroshi Kaneda and Akihiko Tago (Nakano et al. 2007b, 2008b) at an average unfiltered magnitude of 8.7 on 2007 December 25.35 UT (MJD 54,459.35) and 2007 December 26.38 UT (MJD 54,460.38), respectively. Spectroscopic observation of the object by Yamaoka et al. (2007) revealed the presence of several Fe II multiplets, making V459 Vul an Fe II-type nova. A candidate progenitor with $m_{\rm r} \sim 20$ was identified in the red POSS-II plates, but no IR counterpart was found within 2MASS. The progenitor magnitude found suggested a TN-outburst amplitude of $\sim 12.5 \text{ mag}$ in B (Henden & Munari 2008). The maximum magnitude of V459 Vul was $m_V = 7.58$ on 2007 December 27.25 UT (MJD 54,461.25), with t_2 and t_3 times given as 18 ± 2 days, and 30 ± 2 days, respectively (Poggiani 2010), making this a fast nova. Using photometric data, Poggiani (2010) went on to obtain an extinction of $A_V = 2.75 \pm 0.38$, an absolute magnitude range between -8.7and -7.7, and a WD mass in the range of 0.9–1.1 M_{\odot} .

The SMEI light curve is displayed in Figure 8, and unfortunately the initial rise of the nova appears not to be caught by the instrument. Several points between December 25.85 UT and 27.54 UT (MJD 54,459.85-54,461.54) have also been omitted owing to poor r values and a large scatter in the R.A. and decl. of the object early on; as such, any possible PMH has been missed. The peak magnitude of the light curve is found to be $m_{\rm SMEI} = 6.59 \pm 0.04$ on 2007 December 28.11^{+0.11}_{-0.07} UT (MJD 54,462.11). Based on a linear extrapolation between the peak and last data point of the SMEI light curve, the estimated t_2 time is given as ~19.4 days. The SMEI peak is 0.99 mag brighter and 0.86 days later than V-band measurements given in Poggiani (2010), and the estimated t_2 value is slightly larger; these deviations may be due to the difference in bandpass. The peak AAVSO V-band magnitude, however, disagrees with that of Poggiani (2010), giving $m_V = 7.39 \pm 0.01$ on 2007 December 27.98 (MJD 54,461.98) reducing previous discrepancies. In comparison, the peak of the *R*-band AAVSO data is much closer to the SMEI peak as it occurs on 2007 December 27.97 UT (MJD 54,461.97) at $m_R = 6.57 \pm 0.01$. The SMEI light curve also indicates that there may be several oscillations in the decline with amplitudes of order of a few tenths of a magnitude. These oscillations are not evident in the AAVSO data and illustrate how high-cadence observations can reveal new features.

3.7. V2491 Cygni

The discovery of Nova V2491 Cyg ($\alpha = 19^{h}43^{m}01^{\circ}96$, $\delta = +32^{\circ}19'13''8$; J2000) was presented in Nakano et al. (2008a). The nova reached maximum $m_V = 7.45 \pm 0.05$ on 2008 April 11.37 \pm 0.01 UT (MJD 54,567.37; this is the peak of the AAVSO V-band data in Figure 9, Munari et al. 2011). Spectra indicated that it belongs to the He/N class of novae (Helton et al. 2008). V2491 Cyg had a rapid optical decline with a $t_2 = 4.8$ days and as such is classified as a very fast nova (Darnley et al. 2011; Munari et al. 2011). Approximately 15 days after eruption, a secondary maximum was observed with $m_V = 9.49 \pm 0.03$. Based on the spectra of the nova, many authors (e.g., Tomov et al. 2008; Page et al. 2010) believe that the object is in fact an RN. Using the interstellar Na I line, a reddening of $E_{B-V} = 0.23 \pm 0.01$ was derived.

Unfortunately, SMEI was only able to detect the nova at a few points around the nova peak and its initial decline. These data are presented in Figure 9, in which we have allowed data with r < 0.4 for completeness. The maximum magnitude is given as $m_{\rm SMEI} = 7.36 \pm 0.05$ on 2008 April 10.89 \pm 0.04 UT (MJD 54,566.89), slightly earlier and brighter than the *V*-band peak given in Munari et al. (2011), but matching quite well the unfiltered magnitude of 7.7 on 2008 April 10.8 UT (MJD 54,566.8) presented in Nakano et al. (2008a). Owing to a lack of data, we were unable to determine a t_2 time.

THE ASTROPHYSICAL JOURNAL, 820:104 (27pp), 2016 April 1

3.8. QY Muscae

QY Mus ($\alpha = 13^{h}16^{m}36^{s}.44$, $\delta = -67^{\circ}36'47''.8$; J2000) was discovered by Liller et al. (2008), at an unfiltered magnitude of 8.6 on 2008 September 28.998 UT (MJD 54,737.998). The nova then reached a peak unfiltered magnitude of 8.1 on 2008 September 30.397 UT (MJD 54,739.40; green cross within Figure 10). Schwarz et al. (2011) give the t_2 time of this object as approximately 60 days, making this a moderately fast nova.



Figure 9. Optical light curves of V2491 Cyg. SMEI data with a lower r than optimal have been included for completeness. Data from the AAVSO have been plotted for comparison (see Figure 1 for data legend).

The SMEI light curve (see Figure 10) indicates that the nova reached a peak magnitude of $m_{\rm SMEI} = 6.93 \pm 0.04$ on 2008 September 28.63 \pm 0.2 UT (MJD 54,737.63). This magnitude is significantly brighter than both the V-band AAVSO peak and unfiltered peak given in Liller et al. (2008). The variance between the peak magnitudes may be due to differences in the bandpasses of the instruments. However, upon an examination of the location of the nova within the SMEI sky maps, the discrepancy is most likely caused by neighboring bright stars (again, simultaneous fitting of these objects was conducted along with the fitting of additional bright stars in the larger surrounding region). The decay of the SMEI light curve is quite consistent with that from AAVSO data. However, the scatter in the SMEI light curve is evident and is due to problems in fitting the source as it approaches the limiting SMEI magnitude and contamination from the surrounding bright stars. We are unable to give an estimate of the t_2 value of the nova owing to the large scatter in the light curve at this time.

3.9. V5580 Sagittarii

Nova V5580 Sgr ($\alpha = 18^{h}22^{m}01^{s}39$, $\delta = -28^{\circ}02'39''8$; J2000) was discovered by Liller (2008) at approximately 8th magnitude using an "orange" filter on 2008 November 29.04 UT (MJD 54,799.04). The variable was also present on November 23.037 UT (MJD 54,793.037) at a magnitude of approximately 10.3, but was not visible (mag > 11.0) on November 20.035 (MJD 54,790.035). The nova rose to peak on 2008 November 29.999 UT (MJD 54,799.999) with a magnitude of 7.8 (Liller 2008).



Figure 10. Optical light curves of QY Mus. The SMEI data are very noisy owing to contamination from neighboring bright stars. AAVSO data are given for comparison (see Figure 1 for data legend). The green cross represents the peak unfiltered magnitude given by Liller et al. (2008). The inset represents data from around the peak of the SMEI light curve only. These data at maximum are to be treated with caution as they are contaminated by light from neighboring bright stars.

Although the nova is detected by SMEI, only a few reliable points were obtained owing to the passage of Venus and its poor subtraction. For this reason we have decided not to present the light curve. The peak magnitude found for the object is at $m_{\text{SMEI}} = 7.01 \pm 0.04$ on 2008 November 30.85 UT (MJD 54,800.85). This is slightly later and brighter than values mentioned within Liller (2008). Unfortunately, no t_2 time can be derived from the SMEI light curve, and one cannot be found in the literature.

4. DISCUSSION AND CONCLUSION

We present here SMEI archival data for nine known Galactic novae, V1187 Sco, V2467 Cyg, V458 Vul, V597 Pup, V459 Vul, V2491 Cyg, QY Mus, V5580 Sgr, and V5583 Sgr (which are faint with respect to the novae presented in Paper I: RS Oph, V1280 Sco, V598 Pup, and KT Eri, and comparable to T Pyxidis in Hounsell et al. 2012; Surina et al. 2014). Light curves for eight of the eruptions have been displayed, and for all nine we have determined basic observational properties, including the time of peak, maximum magnitude, and where possible the t_2 time (see Table 1). Our work indicates that although these novae possess magnitudes that are at or below the optimal detection limits of SMEI, we are still able to produce light curves that in many cases contain more data at and around the initial rise, peak, and decline than are found within other variable star catalogs such as the AAVSO or the Nova Atlas (the limits here may well be due to sampling and distance).

The majority of novae examined within this paper are of the very fast (5 out of 9) or fast (2 of 9) speed classes rather than the moderately fast, slow, or very slow type. This is not surprising as novae that decline more rapidly are more intrinsically luminous and as such more likely to possess peak magnitudes above the SMEI detection threshold (distance and reddening factors may play a part here as well). QY Mus, however, does belong to the moderately fast speed class, which in turn implies a slow photometric evolution (the slowest nova examined in both this paper and Paper 1), and as such has produced one of the noisiest light curves owing to the faintness of the event (and star crowding issues).

Whenever possible our SMEI light curves have been compared to multi-filter data from the AAVSO and Nova Atlas. In the majority of cases the SMEI data have been seen to match the *R* band (specifically R_{AAVSO}) much better than those of any other filter (unsurprising given that the peak throughput of the instrument is 700 nm). However, for V1187 Sco the SMEI magnitude is best matched with the I-band Nova Atlas data. This discrepancy between favored bandpasses could be caused by contamination from neighboring stars, as well as different filters being used in both the Nova Atlas and AAVSO catalogs, which come from several different telescopes and as such vary in peak throughput. It could also be due to the intrinsic spectral energy distribution of the nova, and/or the column of extinction toward the system. An additional point to note is that V1187 Sco is a fast ONe-type nova and so would have had much stronger line emission early on in the eruption relative to slower, more continuum-dominated novae. The strong H α emission would affect the R and I filters and may well be a more likely explanation for the difference in which one is favored.

Four of the novae examined (V1187 Sco, V458 Vul, V459 Vul, and V5580 Sgr) possess peak SMEI magnitudes that are brighter than their recorded *V*-band maxima and at slightly later times. Again this may simply be due to the

different filters being considered here, which is supported by the fact that two of the novae possess peaks that are closer in time and magnitude to the given *R*-band maxima. On the other hand, four novae (QY Mus, V2467 Cyg, V2491 Cyg, V5583 Sgr) have peak SMEI magnitudes that are earlier and brighter than the values recorded in the literature. For QY Mus we believe that the value of the SMEI maximum is inaccurate owing to severe contamination from neighboring sources. However, for V2467 Cyg, V2491 Cyg, and V5583 Sgr, although the difference in magnitude may be due to the difference in bandpass, the earlier occurrence could suggest that we obtained a more precise time of peak due to bettersampled light curves.

Looking at specific features within the SMEI nova light curves, five of the objects examined possess final rise data (V1187 Sco, V2467 Cyg, V597 Pup, V459 Vul, V5583 Sgr), which for the majority of cases (V5583 Sgr excluded) were not previously observed. Within the final rise of both V597 Pup and V5583 Sgr the presence of a PMH is suggested. Unfortunately, the correlation coefficient (r) of the data for the PMHs is in both cases less than optimal (the first two points of the V597 Pup halt have an r value of 0.39 and 0.35; the V5583 Sgr halt has r values of 0.23, 0.26, and 0.26, respectively), and so their reality is questionable. However, in the case of V5583 Sgr, work by Holdsworth et al. (2014) would suggest that the SMEI PMH is real as they too find this halt using STEREO data. It should also be noted that the duration of each halt observed is consistent with the speed class of the nova and occurs within the Δm_{SMEI} (0.63–1.74 mag) and Δt (0.49–2.19 days) ranges found in Paper I.

Throughout this work and Paper I, we have used the traditional definition of a PMH that occurs 1-2 mag below peak (see, e.g., McLaughlin). However, we should note that until there is a strong physical basis for the PMH, any such definition is relatively arbitrary and, for example, may be dependent on parameters such as the WD mass or accretion rate.

With this in mind, we must re-evaluate both nova V2467 Cyg and V5583 Sgr. Examining the rise of nova V2467 Cyg, there may be evidence for a PMH on 2007 March 16.07 UT (MJD 54,175.07) at an average $m_{\rm SMEI} = 6.56 \pm 0.05$ that lasts ~0.14 days (see inset of Figure 6). This PMH has a Δm_{SMEI} of 0.26 mag, much smaller than that previously observed, and a Δt of 0.49 days. The r values associated with these PMH data are also all above 0.6, making it a very real feature of the light curve. In Section 3.2nova V5583 Sgr was found to have several PMH-like features, and although one seemed to fit our "standard" definition of a PMH, it consisted of data with lower r values than optimal. The third feature seen, however, has data with r > 0.35 (again less than optimal, but better than the earlier features) and occurs on 2009 August 6.75 UT (MJD 55,049.75) at $m_{\rm SMEI} = 7.18 \pm 0.01$ and lasts for 0.21 days. The PMH has an Δm_{SMEI} of 0.25 mag and a Δt of 0.33 days, both of which are much smaller than that previously observed.

During the initial decline from maximum, both V1187 Sco and V459 Vul seem to display oscillations within their light curves. This rapid observed variability can only be appreciated and realized with a high-cadence instrument such as SMEI. Moreover, the unusual nova V458 Vul has several multi flaring events shortly after eruption for which SMEI has again provided a highly detailed light curve. This light curve closely matches that observed in the AAVSO *R* band. With its closure in 2011 September the SMEI archive contains 8.5 yr of all-sky high cadence (102 minutes) data, which is now in a static state. Both Paper I and the present work clearly indicate how important it is to examine the data of all-sky facilities such as SMEI with regard to transient events. As such, we will continue to investigate this archive for other transient and variable stars, as well as for additional novae that may have been missed during their eruptions.

 Table 2

 SMEI Light-curve Data for Nova V1187 Sco

Date-Obs MJD	SMEI Mag	SMEI Error	r
53,215.68	7.47	0.05	0.73
53,217.52	7.49	0.05	0.72
53,217.66	7.62	0.05	0.70
53.217.80	7.59	0.05	0.71
53.218.01	7.67	0.05	0.71
53,218.08	7.73	0.06	0.69
53.218.30	7.66	0.05	0.70
53.218.37	7.78	0.06	0.76
53.218.86	7.87	0.06	0.74
53,219.92	7.37	0.05	0.63
53,219,99	7.17	0.04	0.64
53,220.06	7.29	0.05	0.64
53,220,13	7.26	0.05	0.63
53.220.20	7.07	0.04	0.66
53,220,27	7.20	0.04	0.61
53,220,34	7.02	0.04	0.65
53,220.41	7.03	0.04	0.63
53.220.48	7.03	0.04	0.64
53,220.55	7.01	0.04	0.64
53,220.62	7.12	0.04	0.64
53.220.70	7.04	0.04	0.62
53.220.77	6.95	0.04	0.63
53.220.91	7.09	0.04	0.57
53,220,98	7.06	0.04	0.60
53.221.26	6.97	0.04	0.61
53,221,33	7.11	0.04	0.58
53,221.40	7.15	0.04	0.58
53.221.47	7.20	0.04	0.54
53,221,54	7.29	0.05	0.57
53.221.61	7.32	0.05	0.57
53,221.68	7.33	0.05	0.57
53,221.75	7.43	0.05	0.55
53,221.82	7.29	0.05	0.55
53,221.89	7.25	0.04	0.56
53,221.96	7.52	0.05	0.53
53,222.04	7.54	0.05	0.52
53,222.18	7.50	0.05	0.51
53,222.25	7.46	0.05	0.52
53,222.32	7.53	0.05	0.50
53,222.39	7.63	0.05	0.50
53,222.53	7.41	0.05	0.51
53,222.60	7.54	0.05	0.49
53,222.67	7.63	0.05	0.47
53,222.81	7.53	0.05	0.50
53,222.88	7.60	0.05	0.48
53,222.95	7.52	0.05	0.49
53,223.02	7.49	0.05	0.49
53,223.09	7.53	0.05	0.49
53,223.23	7.49	0.05	0.47
53,223.80	7.70	0.06	0.46
53,223.87	7.57	0.05	0.45
53,223.94	7.81	0.06	0.43
53,224.01	7.60	0.05	0.44
53,224.29	7.94	0.06	0.40

Table 2 (Continued) Date-Obs MJD SMEI Mag SMEI Error r 53.224.36 0.44 7.70 0.05 53,224.43 7.52 0.05 0.4453.224.50 7.91 0.06 0.41 53,224.57 7.84 0.06 0.43 53,224.64 7.74 0.06 0.42 53,224.71 7.81 0.06 0.41 53,224.79 7.85 0.06 0.4053,224.86 7.77 0.06 0.42 53,224.93 7.79 0.06 0.41 53,225.00 7.81 0.06 0.42 53,225.21 7.91 0.06 0.4253,225.28 7.56 0.05 0.48 53,225.35 7.83 0.44 0.06 53,225.42 7.91 0.06 0.42 53.225.49 7.94 0.06 0.42 7.84 0.44 53,226.20 0.06 53,226.27 7.95 0.06 0.43 53,226.34 8.05 0.06 0.42 53,226.48 8.30 0.07 0.39 0.38 53.226.55 8.48 0.08 53,226.62 8.08 0.07 0.38 53,226.83 0.35 8.42 0.08 53,226.90 8.46 0.08 0.35 53,226.97 8.38 0.07 0.33 53,227.04 8.08 0.35 0.07 53,227.26 8.77 0.09 0.31 53,227.33 8.42 0.08 0.36 53.227.40 8.31 0.07 0.34 53 227 54 8.37 0.07 0.33 53,229.45 8.42 0.08 0.33 53,229.94 8.38 0.07 0.36 53,230.01 8.47 0.080.31 53,230.79 8.69 0.09 0.37 53,231.21 9.49 0.12 0.35 53,231.28 9.37 0.12 0.34 53,231.91 9.08 0.10 0.31

R. Hounsell acknowledges support from the Space Telescope Science Institute and the University of Illinois at Urbana– Champaign. The USAF/NASA SMEI is a joint project of the University of California San Diego, Boston College, the University of Birmingham (UK), and the Air Force Research Laboratory. P.P.H., A.B., B.V.J., and J.M.C. acknowledge support from NSF grant ATM-0852246 and NASA grant NNX08AJ11G. A.W.S. acknowledges support from NSF grant AST-1009566. S.S. acknowledges partial support from NASA and NSF grants to ASU. We would like to thank Dr. Zach Cano (Centre for Astrophysics and Cosmology at the University of Iceland) for his thoughtful ideas and help during this project. Finally, we would like to thank the referee for their helpful comments and guidance.

APPENDIX

Tables 2–18 within this appendix contain the MJD, SMEI Mag, and SMEI Error for each nova examined within this paper and Paper I. In addition, r is given for novae presented within this paper only.

 Table 3

 SMEI Light-curve Data for Nova V2467 Cyg

Table 4		
SMEI Light-curve Data for Nova	V458	Vul

5410088 13.07 0.67 0.31 55.312.22 8.57 0.10 0.44 54107.16 10.63 0.22 0.42 55.317.33 8.04 0.07 0.66 541.07.16 10.63 0.22 0.42 55.321.03 8.04 0.07 0.66 541.73.52 10.26 0.19 0.55 55.322.07 8.16 0.08 0.66 54.173.87 8.85 0.11 0.55 55.322.49 8.36 0.09 0.55 54.174.01 8.78 0.06 0.53 55.322.91 8.39 0.09 0.55 54.174.11 7.77 0.06 0.53 54.322.91 8.38 0.09 0.55 54.174.15 7.77 0.06 0.53 54.323.45 8.33 0.09 0.4 54.174.23 7.14 0.04 0.53 54.323.45 8.33 0.09 0.4 54.174.37 7.12 0.04 0.63 54.323.45 8.33 0.09 0.4 </th <th>Date-Obs MJD</th> <th>SMEI Mag</th> <th>SMEI Error</th> <th>r</th> <th>Date-Obs MJD</th> <th>SMEI Mag</th> <th>SMEI Error</th> <th>r</th>	Date-Obs MJD	SMEI Mag	SMEI Error	r	Date-Obs MJD	SMEI Mag	SMEI Error	r
54,161,38 10.30 0.19 0.30 54,321,86 8.06 0.08 0.05 54,172,7 9.10 0.11 0.54 54,322,00 8.15 0.08 0.06 54,173,52 10.26 0.19 0.55 54,322,07 8.16 0.08 0.06 54,173,30 9.14 0.11 0.57 54,322,47 8.36 0.09 0.05 54,173,40 8.85 0.15 0.58 54,322,49 8.36 0.09 0.05 54,174,01 8.78 0.09 0.53 54,322,49 8.36 0.09 0.05 54,174,03 7.09 0.06 0.53 54,323,52 8.38 0.09 0.05 54,174,21 7.77 0.06 0.53 54,323,52 8.33 0.09 0.44 54,174,21 7.17 0.06 0.53,33,33 8.38 0.09 0.44 54,174,40 6.80 0.04 0.53 54,323,55 8.33 0.09 0.44	54,160.88	13.07	0.67	0.31	54,321.22	8.57	0.10	0.40
54,167.16 10.63 0.22 0.42 54,321.93 8.04 0.07 0.06 54,173.27 9.10 0.11 0.53 54,322.07 8.16 0.08 0.66 54,173.80 9.14 0.11 0.57 54,322.35 8.56 0.09 0.55 54,174.01 8.78 0.09 0.55 54,322.37 8.36 0.09 0.55 54,174.01 8.78 0.09 0.53 54,322.91 8.39 0.09 0.55 54,174.15 7.77 0.06 0.52 54,322.91 8.39 0.09 0.55 54,174.27 7.12 0.04 0.58 54,323.48 8.30 0.08 0.44 54,174.37 7.12 0.04 0.63 54,323.48 8.33 0.09 0.44 54,174.37 6.83 0.04 0.63 54,323.44 8.43 0.09 0.44 54,174.57 7.12 0.04 0.63 54,324.44 8.43 0.09 0.44	54,161.38	10.30	0.19	0.30	54,321.86	8.06	0.08	0.62
54,171.27 0.10 0.11 0.54 54,322.00 8.16 0.08 0.06 54,173.50 9.14 0.11 0.57 54,322.14 8.40 0.09 0.65 54,173.80 9.14 0.11 0.57 54,322.34 8.56 0.09 0.55 54,173.41 9.85 0.15 0.58 54,322.39 8.30 0.08 0.55 54,174.01 8.78 0.09 0.55 54,322.99 8.38 0.09 0.55 54,174.03 7.90 0.06 0.52 54,323.12 8.36 0.09 0.55 54,174.22 7.77 0.06 0.58 54,332.12 8.38 0.09 0.44 54,174.30 7.06 0.04 0.67 54,323.55 8.33 0.09 0.44 54,174.37 7.12 0.04 0.63 54,323.42 8.43 0.09 0.44 54,174.37 7.02 0.04 0.62 54,323.41 8.43 0.09 0.44 54,174.41 6.80 0.04 0.62 54,323.41 8.47	54,167.16	10.63	0.22	0.42	54,321.93	8.04	0.07	0.66
54,175.32 10.26 0.19 0.55 54,322.07 8.16 0.08 0.06 54,173.30 9.14 0.11 0.57 54,322.35 8.56 0.09 0.55 54,174.01 8.78 0.09 0.55 54,322.37 8.26 0.09 0.55 54,174.01 8.78 0.09 0.55 54,322.91 8.29 0.09 0.55 54,174.15 7.77 0.06 0.52 54,322.91 8.33 0.09 0.55 54,174.3 7.14 0.04 0.58 54,323.48 8.30 0.08 0.44 54,174.37 7.12 0.04 0.53 54,323.62 8.45 0.09 0.44 54,174.37 7.12 0.04 0.63 54,323.62 8.43 0.09 0.44 54,174.37 6.63 0.04 0.66 54,324.64 8.47 0.09 0.44 54,175.07 6.83 0.04 0.66 54,324.51 8.18 0.08 0.55	54,170.27	9.10	0.11	0.54	54,322.00	8.15	0.08	0.65
54,173.80 9,14 0,11 0.57 54,322,15 8,40 0.09 0.65 54,173.91 9,85 0,15 0.58 54,322,40 8,36 0.09 0.55 54,174.01 8,78 0.09 0.55 54,322,40 8,30 0.09 0.55 54,174.08 7,90 0.06 0.53 54,322,99 8,38 0.09 0.55 54,174.13 7,17 0.06 0.58 54,323,12 8,36 0.09 0.65 54,174.31 7,14 0.04 0.57 54,323,12 8,33 0.09 0.44 54,174.37 7,12 0.04 0.63 54,323,82 8,38 0.09 0.44 54,174.41 6,80 0.04 0.69 54,323,80 8,45 0.09 0.44 54,174.37 7,12 0.04 0.61 54,324,04 8,47 0.09 0.44 54,173.01 6,80 0.04 0.62 54,324,04 8,41 0.09 0.65 54,175.17 6,80 0.04 0.65 54,324,14 8,41	54,173.52	10.26	0.19	0.55	54,322.07	8.16	0.08	0.65
54,173.37 8.85 0.10 0.56 54,322.35 8.56 0.09 0.55 54,174.01 8.78 0.09 0.55 54,322.49 8.30 0.09 0.55 54,174.01 8.78 0.09 0.55 54,322.91 8.30 0.09 0.55 54,174.15 7.77 0.06 0.52 54,323.12 8.33 0.09 0.05 54,174.13 7.14 0.04 0.58 54,323.42 8.43 0.09 0.04 54,174.37 7.12 0.04 0.63 54,323.62 8.45 0.09 0.44 54,174.37 7.12 0.04 0.63 54,323.62 8.45 0.09 0.44 54,174.37 7.12 0.04 0.66 54,323.63 8.38 0.09 0.44 54,174.57 6.80 0.04 0.66 54,324.01 8.47 0.09 0.4 54,174.07 6.80 0.04 0.65 54,324.01 8.47 0.09 0.4 54,174.97 6.80 0.04 0.65 54,324.18 8.20 <	54,173.80	9.14	0.11	0.57	54,322.14	8.40	0.09	0.60
54,173,94 9,85 0.15 0.88 54,322,49 8,36 0.09 0.55 54,174,08 7,90 0.06 0.53 54,322,91 8,39 0.09 0.55 54,174,15 7,77 0.06 0.52 54,322,98 8,38 0.09 0.05 54,174,22 7,77 0.06 0.58 54,323,12 8,36 0.09 0.04 54,174,30 7,16 0.04 0.57 54,323,35 8,33 0.09 0.04 54,174,30 7,16 0.04 0.63 54,323,35 8,38 0.09 0.04 54,174,47 6,80 0.04 0.69 54,323,83 8,38 0.09 0.4 54,174,47 6,80 0.04 0.60 54,324,94 8,47 0.09 0.4 54,175,00 6,64 0.04 0.66 54,324,34 8,48 0.08 0.5 54,175,01 6,48 0.03 0.67 54,324,46 8,41 0.09 0.4 54,175,14 6,55 0.03 0.62 54,324,45 8,51 <td< td=""><td>54,173.87</td><td>8.85</td><td>0.10</td><td>0.56</td><td>54,322.35</td><td>8.56</td><td>0.09</td><td>0.55</td></td<>	54,173.87	8.85	0.10	0.56	54,322.35	8.56	0.09	0.55
54,174,01 8,78 0.09 0.55 54,322,70 8,20 0.08 0.55 54,174,15 7,77 0.06 0.53 54,322,91 8,39 0.09 0.55 54,174,15 7,77 0.06 0.55 54,323,12 8,36 0.09 0.45 54,174,33 7,14 0.04 0.55 54,323,42 8,33 0.09 0.44 54,174,43 7,12 0.04 0.63 54,323,62 8,45 0.09 0.44 54,174,47 7,12 0.04 0.63 54,323,83 8,38 0.09 0.44 54,174,77 6,80 0.04 0.65 54,323,90 8,45 0.09 0.44 54,174,77 6,80 0.04 0.66 54,324,04 8,47 0.09 0.44 54,174,77 6,80 0.04 0.65 54,324,18 8,20 0.08 0.55 54,175,07 6,48 0.03 0.67 54,324,45 8,41 0.09 0.4 54,175,14 6,55 0.03 0.65 54,324,51 8,41	54,173.94	9.85	0.15	0.58	54,322.49	8.36	0.09	0.53
54,174,18 7,90 0.06 0.53 54,322.91 8.39 0.09 0.5 54,174,12 7,77 0.06 0.52 54,322.98 8.38 0.09 0.5 54,174,30 7,14 0.04 0.52 54,323.12 8.36 0.09 0.4 54,174,37 7,12 0.04 0.57 54,323.55 8.33 0.09 0.4 54,174,47 7,12 0.04 0.63 54,323.55 8.38 0.09 0.4 54,174,70 6.80 0.04 0.60 54,323.40 8.47 0.09 0.4 54,174,70 6.80 0.04 0.61 54,324.11 8.56 0.09 0.4 54,175,00 6.65 0.04 0.65 54,324.15 8.18 0.08 0.5 54,175,00 6.65 0.04 0.65 54,324.45 8.41 0.09 0.4 54,175,10 6.61 0.04 0.63 54,324.46 8.41 0.09 0.4 54,175,28 6.60 0.03 0.57 54,324.54 8.50 0.09<	54,174.01	8.78	0.09	0.55	54,322.70	8.20	0.08	0.58
54,174,12 7,77 0.06 0.52 54,322,12 8,38 0.09 0.5 54,174,32 7,17 0.06 0.58 54,323,12 8,36 0.08 0.44 54,174,30 7,06 0.04 0.57 54,323,45 8,33 0.09 0.44 54,174,37 7,12 0.04 0.63 54,323,62 8,45 0.09 0.44 54,174,71 6,83 0.04 0.60 54,323,62 8,45 0.09 0.44 54,174,71 6,83 0.04 0.60 54,324,04 8,47 0.09 0.44 54,174,72 6,80 0.04 0.66 54,324,11 8,56 0.09 0.44 54,175,07 6,48 0.04 0.66 54,324,11 8,50 0.09 0.45 54,175,17 6,61 0.04 0.65 54,324,45 8,81 0.08 0.5 54,175,12 6,61 0.04 0.65 54,324,61 8,47 0.09 0.44 54,175,42 6,61 0.04 0.55 54,324,61 8,47 <	54,174.08	7.90	0.06	0.53	54,322.91	8.39	0.09	0.52
54,174.22 7,77 0.06 0.58 54,323.12 8.36 0.09 0.5 54,174.30 7,16 0.04 0.57 54,323.55 8.33 0.09 0.44 54,174.57 7,12 0.04 0.63 54,323.55 8.33 0.09 0.44 54,174.47 6.80 0.04 0.69 54,323.53 8.38 0.09 0.44 54,174.71 6.83 0.04 0.62 54,323.04 8.45 0.09 0.44 54,174.71 6.83 0.04 0.66 54,324.04 8.47 0.09 0.44 54,175.07 6.48 0.04 0.68 54,324.18 8.20 0.08 0.55 54,175.07 6.48 0.03 0.67 54,324.46 8.41 0.09 0.45 54,175.14 6.51 0.04 0.65 54,324.45 8.50 0.09 0.44 54,175.31 6.61 0.04 0.65 54,324.51 8.11 0.08 0.5 54,175.53 6.30 0.03 0.57 54,324.52 8.31 <	54,174.15	7.77	0.06	0.52	54,322.98	8.38	0.09	0.53
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,174.22	7.77	0.06	0.58	54,323.12	8.36	0.09	0.51
54,174,50 7,06 0.04 0.57 54,323,55 8,43 0.09 0.4 54,174,57 7,12 0.04 0.63 54,323,62 8,45 0.09 0.4 54,174,71 6,83 0.04 0.60 54,323,83 8,38 0.09 0.4 54,174,79 6,80 0.04 0.61 54,323,40 8,47 0.09 0.4 54,174,79 6,80 0.04 0.61 54,324,18 8,20 0.08 0.5 54,175,00 6,65 0.04 0.65 54,324,41 8,40 0.09 0.4 54,175,00 6,65 0.04 0.65 54,324,41 8,40 0.09 0.4 54,175,14 6,55 0.03 0.62 54,324,61 8,47 0.09 0.4 54,175,14 6,55 0.03 0.67 54,324,41 8,51 0.09 0.4 54,175,42 6,61 0.04 0.55 54,324,47 8,31 0.08 0.5 54,175,45 6,50 0.03 0.67 54,324,81 8,11 0.08<	54,174.43	7.14	0.04	0.58	54,323.48	8.30	0.08	0.49
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,174.50	7.06	0.04	0.57	54,323.55	8.33	0.09	0.49
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,174.57	7.12	0.04	0.63	54,323.62	8.45	0.09	0.44
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,174.64	6.80	0.04	0.59	54,323.83	8.38	0.09	0.49
54,174.79 6.80 0.04 0.62 54,324.04 8.47 0.09 0.44 54,174.86 6.75 0.04 0.68 54,324.18 8.20 0.08 0.55 54,175.00 6.65 0.04 0.65 54,324.25 8.18 0.08 0.55 54,175.07 6.48 0.03 0.67 54,324.46 8.41 0.09 0.4 54,175.14 6.55 0.03 0.62 54,324.46 8.47 0.09 0.4 54,175.14 6.61 0.04 0.63 54,324.46 8.47 0.09 0.4 54,175.28 6.60 0.03 0.55 54,324.28 8.51 0.09 0.4 54,175.56 6.30 0.03 0.67 54,324.96 8.23 0.08 0.5 54,175.56 6.30 0.03 0.67 54,324.91 8.10 0.08 0.5 54,176.52 7.20 0.05 0.32 54,323.10 8.17 0.08 0.5 54,176.62 7.46 0.05 0.32 54,325.45 8.17 0.	54,174.71	6.83	0.04	0.60	54,323.90	8.45	0.09	0.47
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,174.79	6.80	0.04	0.62	54,324.04	8.47	0.09	0.46
54,174.93 6.64 0.04 0.68 54,324.18 8.20 0.08 0.55 54,175.00 6.65 0.04 0.65 54,324.25 8.18 0.09 0.55 54,175.07 6.48 0.03 0.67 54,324.24 8.50 0.09 0.44 54,175.14 6.51 0.04 0.63 54,324.46 8.47 0.09 0.44 54,175.28 6.60 0.03 0.55 54,324.45 8.31 0.08 0.55 54,175.49 6.50 0.03 0.67 54,324.42 8.51 0.09 0.44 54,175.65 6.30 0.03 0.67 54,324.42 8.23 0.08 0.55 54,175.65 6.30 0.03 0.67 54,325.10 8.17 0.08 0.55 54,176.62 7.46 0.05 0.32 54,325.24 8.29 0.08 0.55 54,176.62 7.46 0.05 0.37 54,325.45 8.17 0.08 0.65 54,176.62 7.46 0.05 0.37 54,325.45 8.17	54,174.86	6.75	0.04	0.61	54,324.11	8.56	0.09	0.49
54,175.00 6.65 0.04 0.65 $54,324,25$ $8,18$ 0.08 0.55 $54,175.07$ 6.48 0.03 0.67 $54,324,46$ $8,41$ 0.09 0.44 $54,175.14$ 6.55 0.03 0.62 $54,324,46$ $8,47$ 0.09 0.44 $54,175.21$ 6.61 0.04 0.63 $54,324,57$ $8,31$ 0.08 0.55 $54,175.42$ 6.61 0.04 0.55 $54,324,75$ $8,31$ 0.09 0.44 $54,175.42$ 6.61 0.04 0.56 $54,324,52$ $8,51$ 0.09 0.44 $54,175.45$ 6.30 0.03 0.70 $54,325,10$ $8,17$ 0.08 0.55 $54,175.65$ 6.30 0.03 0.67 $54,325,17$ $8,14$ 0.08 0.55 $54,176.55$ 7.20 0.05 0.32 $54,325,24$ 8.29 0.08 0.55 $54,176.69$ 7.20 0.05 0.37 $54,325,31$ $8,11$ 0.08 0.55 $54,176.67$ 7.15 0.05 0.37 $54,325,52$ $8,13$ 0.08 0.65 $54,176.69$ 7.20 0.05 0.35 $54,325,52$ $8,13$ 0.08 0.66 $54,176.69$ 7.20 0.05 0.35 $54,325,59$ 8.07 0.08 0.66 $54,176.87$ 7.15 0.05 0.35 $54,325,59$ 8.07 0.08 0.66 $54,182.76$ 9.29 0.12 0.43 $54,325,59$ <t< td=""><td>54,174.93</td><td>6.64</td><td>0.04</td><td>0.68</td><td>54,324.18</td><td>8.20</td><td>0.08</td><td>0.56</td></t<>	54,174.93	6.64	0.04	0.68	54,324.18	8.20	0.08	0.56
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.00	6.65	0.04	0.65	54,324.25	8.18	0.08	0.54
54,175,14 6.55 0.03 0.62 $54,324,54$ 8.50 0.09 0.4 $54,175,21$ 6.61 0.04 0.63 $54,324,51$ 8.47 0.09 0.4 $54,175,28$ 6.60 0.03 0.55 $54,324,75$ 8.31 0.08 0.55 $54,175,42$ 6.61 0.04 0.65 $54,324,82$ 8.51 0.09 0.4 $54,175,56$ 6.30 0.03 0.67 $54,324,96$ 8.23 0.08 0.55 $54,175,56$ 6.30 0.03 0.67 $54,325,17$ 8.14 0.08 0.55 $54,175,56$ 6.40 0.05 0.35 $54,325,17$ 8.14 0.08 0.55 $54,176,48$ 7.21 0.05 0.32 $54,325,31$ 8.15 0.08 0.55 $54,176,69$ 7.20 0.05 0.37 $54,325,32$ 8.11 0.08 0.65 $54,176,69$ 7.20 0.05 0.37 $54,325,52$ 8.13 0.08 0.66 $54,176,83$ 7.22 0.05 0.35 $54,325,59$ 8.07 0.08 0.66 $54,176,90$ 7.27 0.05 0.35 $54,325,59$ 8.07 0.08 0.66 $54,182,26$ 9.29 0.12 0.43 $54,325,95$ 8.07 0.08 0.66 $54,182,26$ 9.29 0.12 0.43 $54,325,95$ 8.07 0.08 0.66 $54,182,62$ 9.86 0.15 0.44 $54,326,02$	54,175.07	6.48	0.03	0.67	54,324.46	8.41	0.09	0.50
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.14	6.55	0.03	0.62	54,324.54	8.50	0.09	0.47
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.21	6.61	0.04	0.63	54,324.61	8.47	0.09	0.46
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.28	6.60	0.03	0.55	54,324.75	8.31	0.08	0.51
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.42	6.61	0.04	0.56	54,324.82	8.51	0.09	0.47
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.49	6.50	0.03	0.67	54,324.96	8.23	0.08	0.53
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,175.56	6.30	0.03	0.70	54,325.10	8.17	0.08	0.57
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,175.63	6.40	0.03	0.67	54,325.17	8.14	0.08	0.58
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,176.48	7.21	0.05	0.35	54,325.24	8.29	0.08	0.54
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,176.55	7.20	0.05	0.39	54,325.31	8.15	0.08	0.58
34,176,69 1.20 0.05 0.37 $54,325,45$ 8.17 0.08 0.65 $54,176,76$ 7.15 0.05 0.37 $54,325,52$ 8.13 0.08 0.66 $54,176,69$ 7.22 0.05 0.35 $54,325,52$ 8.07 0.08 0.66 $54,176,90$ 7.27 0.05 0.35 $54,325,73$ 7.98 0.07 0.66 $54,181,77$ 9.32 0.12 0.38 $54,325,73$ 7.98 0.07 0.7 $54,182,48$ 9.18 0.11 0.37 $54,325,80$ 8.00 0.07 0.7 $54,182,55$ 9.60 0.14 0.42 $54,325,95$ 8.07 0.08 0.66 $54,182,62$ 9.86 0.15 0.44 $54,326,02$ 8.13 0.08 0.66 $54,182,69$ 8.55 0.08 0.43 $54,326,02$ 8.14 0.08 0.66 $54,182,69$ 8.55 0.08 0.43 $54,326,09$ 8.14 0.09 0.66 $54,182,69$ 8.55 0.08 0.43 $54,326,53$ 8.34 0.09 0.55 $54,182,83$ 8.89 0.10 0.46 $54,326,55$ 8.32 0.09 0.55 $54,182,90$ 8.44 0.08 0.43 $54,326,55$ 8.32 0.09 0.55 $54,183,46$ 10.53 0.21 0.38 $54,326,55$ 8.32 0.09 0.55 $54,183,46$ 10.53 0.21 0.48 $54,327,57$ <t< td=""><td>54,176.62</td><td>7.46</td><td>0.05</td><td>0.32</td><td>54,325.38</td><td>8.11</td><td>0.08</td><td>0.56</td></t<>	54,176.62	7.46	0.05	0.32	54,325.38	8.11	0.08	0.56
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,176.69	7.20	0.05	0.37	54,325.45	8.17	0.08	0.58
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,176.76	7.15	0.05	0.37	54,325.52	8.13	0.08	0.64
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,176.85	7.22	0.05	0.35	54,325.59	8.07	0.08	0.04
34,161,17 9.52 0.12 0.53 $54,252,15$ 1.96 0.07 0.07 $54,182.26$ 9.29 0.12 0.43 $54,325,80$ 8.00 0.07 0.66 $54,182.26$ 9.18 0.11 0.37 $54,325,88$ 7.98 0.07 0.7 $54,182.55$ 9.60 0.14 0.42 $54,325,95$ 8.07 0.08 0.66 $54,182.62$ 9.86 0.15 0.44 $54,326,02$ 8.13 0.08 0.66 $54,182.69$ 8.55 0.08 0.43 $54,326,02$ 8.13 0.09 0.66 $54,182.76$ 9.38 0.12 0.45 $54,326,23$ 8.34 0.09 0.66 $54,182.76$ 9.38 0.12 0.45 $54,326,23$ 8.34 0.09 0.66 $54,182.76$ 9.38 0.12 0.45 $54,326,30$ 8.43 0.09 0.55 $54,182.76$ 9.38 0.12 0.45 $54,326,58$ 8.26 0.08 0.55 $54,182.76$ 9.38 0.10 0.46 $54,326,58$ 8.26 0.08 0.55 $54,183.46$ 10.53 0.21 0.38 $54,326,58$ 8.63 0.10 0.44 $54,183.61$ 10.44 0.20 0.40 $54,326,66$ 8.63 0.10 0.44 $54,183.75$ 8.78 0.09 0.50 $54,327,43$ 8.51 0.09 0.44 $54,184.03$ 7.73 0.06 0.33 $54,328,06$ <	54,170.90	1.27	0.05	0.35	54,323.00	7.94	0.07	0.08
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,101.77	9.52	0.12	0.38	54,525.75	7.98	0.07	0.70
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,182.20	9.29	0.12	0.45	54,525.60	8.00 7.09	0.07	0.07
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,182.48 54,182.55	9.18	0.11	0.37	54,525.00	7.98	0.07	0.71
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54,182.55	9.00	0.14	0.42	54,325.95	8.07	0.08	0.09
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,182.02	9.80	0.15	0.44	54,320.02	8.13	0.08	0.07
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54,182.09	0.35	0.08	0.45	54,326.03	8.14	0.08	0.07
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54 182 83	8.89	0.12	0.45	54 326 30	8.43	0.09	0.01
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54 182 90	8 44	0.08	0.43	54 326 58	8.26	0.09	0.55
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54 183 46	10.53	0.00	0.45	54 326 65	8.20	0.00	0.54
$\begin{array}{c ccccccccccccccccccccccccccccccccccc$	54 183 54	9.43	0.13	0.36	54 326 72	8.47	0.09	0.52
$\begin{array}{cccccccccccccccccccccccccccccccccccc$	54 183 61	10.44	0.20	0.10	54 326 86	8.63	0.10	0.10
$ \begin{array}{cccccccccccccccccccccccccccccccccccc$	54 183 75	8 78	0.09	0.40	54 327 43	8 51	0.09	0.40
54,184.03 7.73 0.06 0.33 54,328.06 8.87 0.11 0.42 54,184.03 7.71 0.06 0.43 54,328.06 8.87 0.11 0.42 54,184.10 7.71 0.06 0.43 54,328.13 8.64 0.10 0.44 54,184.17 7.93 0.06 0.38 54,328.20 8.31 0.08 0.44 54,184.59 7.66 0.06 0.31 54,328.34 8.57 0.10 0.44 54,184.73 8.19 0.07 0.38 54,328.41 8.44 0.09 0.55 54,184.80 7.86 0.06 0.33 54,328.63 8.84 0.11 0.44	54 183 82	8.97	0.10	0.30	54 327 57	8 57	0.10	0.13
54,184.10 7.71 0.06 0.43 54,328.13 8.64 0.10 0.44 54,184.17 7.93 0.06 0.38 54,328.20 8.31 0.08 0.44 54,184.17 7.93 0.06 0.38 54,328.34 8.57 0.10 0.44 54,184.59 7.66 0.06 0.31 54,328.34 8.57 0.10 0.44 54,184.73 8.19 0.07 0.38 54,328.41 8.44 0.09 0.55 54,184.80 7.86 0.06 0.33 54,328.63 8.84 0.11 0.44	54.184.03	7.73	0.06	0.33	54.328.06	8,87	0.11	0.42
54,184.17 7.93 0.06 0.38 54,328.20 8.31 0.08 0.4' 54,184.59 7.66 0.06 0.31 54,328.34 8.57 0.10 0.4' 54,184.73 8.19 0.07 0.38 54,328.41 8.44 0.09 0.5' 54,184.80 7.86 0.06 0.33 54,328.63 8.84 0.11 0.4'	54,184,10	7.71	0.06	0.43	54.328.13	8.64	0.10	0.44
54,184.59 7.66 0.06 0.31 54,328.34 8.57 0.10 0.4' 54,184.73 8.19 0.07 0.38 54,328.41 8.44 0.09 0.5' 54,184.80 7.86 0.06 0.33 54,328.63 8.84 0.11 0.4' 54,184.88 8.17 0.07 0.37 54,328.63 8.84 0.11 0.4'	54.184.17	7.93	0.06	0.38	54.328.20	8.31	0.08	0.47
54,184.73 8.19 0.07 0.38 54,328.41 8.44 0.09 0.55 54,184.80 7.86 0.06 0.33 54,328.63 8.84 0.11 0.41 54,184.88 8.17 0.07 0.37 54,328.63 8.84 0.11 0.41	54,184,59	7,66	0.06	0.31	54.328.34	8.57	0.10	0.47
54,184.80 7.86 0.06 0.33 54,328.63 8.84 0.11 0.44 54,184.88 8.17 0.07 0.37 54,328.70 8.87 0.11 0.44	54.184.73	8.19	0.07	0.38	54.328.41	8.44	0.09	0.52
	54,184.80	7.86	0.06	0.33	54,328.63	8.84	0.11	0.48
$J_{10} J_{10} $	54,184,88	8.17	0.07	0.37	54.328.70	8.87	0.11	0.43
54,184.95 7.81 0.06 0.33	54,184.95	7.81	0.06	0.33				

Date-Obs MJD 54,407.04 54,416.43 54,416.50 54,417.34 54,417.48 54,417.62 54,417.70 54,417.84 54,417.91 54,417.98 54,418.54 54,418.61 54,418.68 54,418.75 54,418.82 54,418.89 54,418.96 54,419.04 54,419.11 54,419.18 54,419.25 54,419.32 54,419.39 54,419.53 54,419.60 54,419.67 54,419.74 54,419.81 54,419.88 54,419.95 54,420.23 54,420.30 54,420.45 54,420.52 54,420.66 54,420.94 54,421.01 54,421.22 54,421.29 54,421.36 54,421.65 54,421.72 54,421.79 54,422.00 54,422.07 54,422.28 54,422.42 54,422.56 54,422.63 54,422.78 54,422.85 54,422.92 54,422.99 54,423.06 54,423.20 54,423.34 54,423.48 54,423.76 54,424.12 54,424.19 54,424.33 54,424.96 54,426.73 54,426.80

Table 5 S

Table 5 SMEI Light-curve Data for Nova V597 Pup				Table 5 (Continued)					
	SMEI Mag	SMEI Error	r	Date-Obs MJD	SMEI Mag	SMEI Error	r		
	9.57	0.13	0.34	54 426 87	8.98	0.10	0.45		
	10.48	0.20	0.43	54.427.36	9.43	0.12	0.52		
	10.30	0.18	0.31	54 427 43	10.06	0.12	0.32		
	9.56	0.13	0.40	54 427 64	10.35	0.19	0.47		
	8.79	0.09	0.37	54 427 71	10.20	0.17	0.45		
	8.34	0.07	0.39	54 427 86	10.73	0.22	0.44		
	7.97	0.06	0.35	54 428 07	11.35	0.29	0.30		
	7.89	0.06	0.40	54 428 14	9 4 9	0.13	0.40		
	7.76	0.06	0.41	54 428 35	9.15	0.13	0.46		
	7.69	0.06	0.41	54 428 42	9.54	0.13	0.47		
	7.02	0.04	0.61	54 428 70	9.43	0.12	0.46		
	6.96	0.04	0.64	54.428.77	9.32	0.12	0.46		
	6.91	0.04	0.66	54.428.84	9.00	0.10	0.46		
	7.03	0.04	0.63	54,428,91	10.19	0.17	0.49		
	7.17	0.04	0.55	54,429,41	12.00	0.40	0.48		
	7.23	0.05	0.55	54,429,62	11.19	0.27	0.44		
	7.24	0.05	0.58	54,429,97	10.49	0.20	0.48		
	7.17	0.04	0.59	54,430.25	11.65	0.34	0.44		
	7.22	0.04	0.56	54,430.61	11.13	0.27	0.48		
	7.33	0.05	0.58	54,430.82	10.74	0.22	0.50		
	7.33	0.05	0.55	54,430,89	11.66	0.34	0.52		
	7.42	0.05	0.52	54431.03	11.25	0.28	0.47		
	7.51	0.05	0.52	54,433.00	10.73	0.22	0.45		
	7.73	0.06	0.49	54.433.22	9.78	0.14	0.52		
	7.76	0.06	0.45	54,434.91	12.96	0.61	0.52		
	8.05	0.07	0.44	54,435,12	9.82	0.15	0.47		
	7.87	0.06	0.47	54,435.19	11.63	0.33	0.50		
	7.93	0.06	0.45	54,435,41	11.07	0.26	0.50		
	7.75	0.06	0.44	54,436.11	11.71	0.35	0.50		
	7.80	0.06	0.44	54,436.89	11.87	0.37	0.44		
	8.28	0.07	0.47	54,437.24	11.02	0.25	0.44		
	8.70	0.09	0.47	54,437.80	11.36	0.30	0.50		
	8.45	0.08	0.44	54,438,58	10.46	0.20	0.47		
	8.59	0.08	0.46	54,438.65	10.05	0.16	0.42		
	8.30	0.07	0.49	54,439.85	11.55	0.32	0.49		
	8.50	0.08	0.46	54,439.92	12.96	0.61	0.49		
	8.95	0.10	0.42	54,441.12	10.28	0.18	0.40		
	8.96	0.10	0.47	54,441.54	11.81	0.36	0.45		
	8.78	0.09	0.45	54,441.61	10.86	0.24	0.46		
	8.72	0.09	0.46	54,441.82	9.37	0.12	0.46		
	8.69	0.09	0.50	54,442.04	10.84	0.23	0.38		
	8.75	0.09	0.48	54,442.18	10.28	0.18	0.38		
	8.72	0.09	0.49	54,442.67	11.42	0.30	0.48		
	9.10	0.11	0.43	54,442.74	9.50	0.13	0.44		
	8.87	0.10	0.55	54,442.88	12.91	0.60	0.33		
	9.38	0.12	0.44	54,443.09	10.21	0.18	0.37		
	9.22	0.11	0.48	54,443.66	10.35	0.19	0.48		
	8.99	0.10	0.50	54,443.87	9.69	0.14	0.37		
	8.86	0.09	0.50	54,444.08	9.98	0.16	0.35		
	9.07	0.10	0.51	54,444.36	10.13	0.17	0.43		
	9.31	0.12	0.50	54,444.58	10.44	0.19	0.34		
	8.97	0.10	0.51	54,444.72	9.67	0.14	0.35		
	9.12	0.11	0.45	54,444.86	10.52	0.20	0.45		
	8.98	0.10	0.46	54,445.07	11.79	0.36	0.44		
	8.70	0.09	0.45	54,445.14	10.24	0.18	0.41		
	9.61	0.13	0.55	54,445.21	10.78	0.23	0.35		
	10.89	0.24	0.54	54,445.92	9.39	0.12	0.34		
	9.82	0.15	0.50	54,446.41	10.21	0.18	0.42		
	10.29	0.18	0.48	54,446.62	9.49	0.13	0.42		
	9.32	0.12	0.43	54,446.98	10.21	0.18	0.47		
	9.54	0.13	0.52	54,447.05	10.10	0.17	0.39		
	8.81	0.09	0.48	54,447.19	9.53	0.13	0.45		
	10.38	0.19	0.49	54,447.26	9.74	0.14	0.45		
	9.00	0.10	0.51	54,447.33	9.31	0.12	0.41		
				,		~ ~ ~ ~	···-		

Table 5(Continued)

Date-Obs MJD	SMEI Mag	SMEI Error	r
54,447.54	9.01	0.10	0.40
54,447.68	9.13	0.11	0.40

 Table 6

 SMEI Light-curve Data for Nova V597 Pup Continued

Date-Obs MJD	SMEI Mag	SMEI Error	r
54,447.75	9.06	0.10	0.40
54,447.82	9.72	0.14	0.39
54,448.04	9.43	0.12	0.35
54,448.11	9.71	0.14	0.42
54,448.25	8.95	0.10	0.49
54,448.46	9.43	0.12	0.39
54,449.02	9.21	0.11	0.39
54,449.52	9.44	0.12	0.37
54,449.59	9.79	0.14	0.41
54,449.66	9.45	0.12	0.40
54,449.87	9.38	0.12	0.40
54,450.01	9.88	0.15	0.39
54,450.15	9.73	0.14	0.49
54,450.36	9.48	0.13	0.42
54,450.50	10.29	0.18	0.45
54,450.57	10.01	0.16	0.44
54,450.64	11.04	0.26	0.42
54,450.86	10.09	0.17	0.54

Table 7SMEI Light-curve Data for Nova V459 Vul

Date-Obs MJD	SMEI Mag	SMEI Error	r
54,431.27	9.79	0.15	0.45
54,431.91	9.02	0.11	0.41
54,449.05	8.18	0.07	0.41
54,459.85	7.06	0.04	0.49
54,459.99	6.86	0.04	0.48
54,461.54	6.88	0.04	0.47
54,461.61	6.85	0.04	0.47
54,461.68	6.79	0.04	0.52
54,461.89	6.65	0.04	0.56
54,461.96	6.65	0.04	0.56
54,462.11	6.59	0.04	0.58
54,462.32	6.86	0.04	0.49
54,462.39	6.88	0.04	0.49
54,462.46	7.52	0.05	0.44
54,462.53	6.93	0.04	0.53
54,462.67	7.11	0.04	0.43
54,462.74	7.00	0.04	0.43
54,462.88	6.89	0.04	0.48
54,462.95	6.99	0.04	0.46
54,463.02	6.88	0.04	0.53
54,463.16	6.96	0.04	0.48
54,463.23	6.87	0.04	0.52
54,463.45	6.84	0.04	0.50
54,463.52	6.82	0.04	0.51
54,463.73	6.82	0.04	0.57
54,463.94	6.94	0.04	0.50
54,464.08	6.95	0.04	0.50
54,464.29	7.05	0.04	0.50
54,464.43	7.07	0.04	0.54
54.464.50	7.11	0.04	0.52

(Continued) Date-Obs MJD SMEI Mag SMEI Error r 54,464.57 0.51 7.19 0.05 54,464.64 7.13 0.04 0.49 54,464.79 7.18 0.05 0.51 54,464.93 7.46 0.05 0.47 54,465.21 7.15 0.05 0.45 54,465.28 7.14 0.04 0.47 0.50 54,465.35 7.14 0.04 54,466.20 7.12 0.04 0.47 54,466.34 7.06 0.04 0.52 54,466.41 7.15 0.05 0.48 54,466.55 7.09 0.04 0.44 54,466.90 7.10 0.04 0.49 54,466.97 7.14 0.05 0.54 54,467.40 7.36 0.05 0.46 0.50 54,467.54 7.25 0.05 54,470.79 7.27 0.48 0.05 54,477.70 8.20 0.07 0.52

 Table 8

 SMEI Light-curve Data for Nova V2491 Cyg

Date-Obs MJD	SMEI Mag	SMEI Error	r
54,566.61	7.46	0.05	0.43
54,566.68	7.46	0.05	0.45
54,566.75	7.42	0.05	0.45
54,566.82	7.37	0.05	0.44
54,566.89	7.36	0.05	0.41
54,566.96	7.54	0.05	0.46
54,567.03	7.56	0.05	0.45
54,567.10	7.89	0.06	0.42
54,567.17	7.90	0.06	0.40
54,567.24	7.63	0.05	0.49
54,567.59	8.64	0.09	0.44
54,567.66	10.67	0.22	0.31
54,567.88	12.85	0.58	0.40

 Table 9

 SMEI Light-curve Data for Nova QYMus

Date-Obs MJD	SMEI Mag	SMEI Error	r
54,705.26	7.89	0.06	0.97
54,705.40	7.89	0.06	0.97
54,705.47	7.56	0.05	0.97
54,705.54	7.87	0.06	0.97
54,705.61	8.33	0.08	0.97
54,705.68	7.75	0.06	0.97
54,708.01	7.83	0.06	0.96
54,708.36	7.95	0.07	0.97
54,708.79	7.86	0.06	0.94
54,709.63	8.77	0.09	0.96
54,710.05	8.12	0.07	0.97
54,710.20	7.81	0.06	0.97
54,710.27	7.66	0.06	0.97
54,710.62	8.29	0.08	0.97
54,713.09	8.05	0.07	0.97
54,713.16	7.92	0.06	0.97
54,713.44	7.92	0.06	0.97

Table 9 Continued)

r 0.97 0.97 0.97 0.97 0.97 0.95 0.96 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.95 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.97 0.96 0.97

(Continued)				(Continued)			
Date-Obs MJD	SMEI Mag	SMEI Error	r	Date-Obs MJD	SMEI Mag	SMEI Error	
54,713.51	8.02	0.07	0.97	54,720.42	8.26	0.08	
54,713.58	8.18	0.07	0.97	54,720.49	8.12	0.07	
54,713.65	7.82	0.06	0.97	54,720.56	7.77	0.06	
54,713.72	7.93	0.06	0.97	54,720.63	8.32	0.08	
54,714.36	7.93	0.06	0.98	54,720.70	8.17	0.07	
54,714.43	7.67	0.06	0.97	54,720.78	7.51	0.05	
54,714.50	7.58	0.06	0.97	54,720.99	7.65	0.06	
54,714.57	7.36	0.05	0.97	54,721.06	7.80	0.06	
54,714.64	7.57	0.05	0.97	54,721.13	8.05	0.07	
54,714.71	7.84	0.06	0.97	54,721.20	7.63	0.06	
54,715.06	7.32	0.05	0.97	54,721.27	7.51	0.05	
54,715,20	7.32	0.05	0.97	54,721,34	7.89	0.06	
54.715.27	7.61	0.06	0.97	54.721.41	8.20	0.07	
54,715,34	7.27	0.05	0.97	54,721,48	7.83	0.06	
54.715.41	7.59	0.06	0.97	54.721.62	8.15	0.07	
54 715 48	7 27	0.05	0.97	54 721 69	8 40	0.08	
54 715 55	7.27	0.05	0.97	54 721 76	8 54	0.09	
54 715 62	7.40	0.05	0.97	54,722.05	8 24	0.07	
54 715 70	7.33	0.05	0.97	54 722.03	8.60	0.07	
54 715 77	7.45	0.05	0.97	54,722.12	8.00	0.09	
54,715.77	7.50	0.05	0.90	54,722.19	8.20	0.07	
54,716.05	7.43	0.05	0.97	54,722.55	0.00 9.42	0.10	
54,716.12	1.01	0.00	0.97	54,722.40	0.43	0.08	
54,710.19	/.//	0.00	0.97	54,722.47	8.00	0.07	
54,/16.26	8.18	0.07	0.97	54,722.54	8.09	0.07	
54,716.33	7.81	0.06	0.97	54,722.61	7.89	0.06	
54,716.40	7.74	0.06	0.97	54,722.68	8.97	0.10	
54,716.47	7.83	0.06	0.97	54,722.75	8.28	0.08	
54,716.54	7.46	0.05	0.97	54,723.03	8.03	0.07	
54,716.61	7.98	0.07	0.97	54,723.10	8.14	0.07	
54,716.68	7.70	0.06	0.97	54,723.18	8.49	0.08	
54,716.75	7.81	0.06	0.97	54,723.32	8.17	0.07	
54,717.04	7.43	0.05	0.97	54,723.39	9.98	0.16	
54,717.11	7.37	0.05	0.97	54,723.53	7.87	0.06	
54,717.18	7.28	0.05	0.97	54,723.60	7.90	0.06	
54,717.25	7.48	0.05	0.97	54,723.67	8.29	0.08	
54,717.32	7.28	0.05	0.97	54,723.74	8.33	0.08	
54,717.39	7.61	0.06	0.98	54,724.16	7.47	0.05	
54,717.46	7.51	0.05	0.97	54,724.23	7.56	0.05	
54,717.53	7.39	0.05	0.97	54,724.30	7.63	0.06	
54,717.60	7.29	0.05	0.97	54,724.45	7.92	0.06	
54,717.74	7.55	0.05	0.96	54,724.52	7.63	0.06	
54,718.02	7.96	0.07	0.97	54,724.59	7.90	0.06	
54,718.16	7.63	0.06	0.97	54,725.08	7.99	0.07	
54,718.23	7.62	0.06	0.97	54,725.22	7.97	0.07	
54,718.38	7.45	0.05	0.97	54,725.29	7.76	0.06	
54,718.45	7.66	0.06	0.97	54,725.50	8.54	0.09	
54,718.52	7.50	0.05	0.97	54,727.76	7.28	0.05	
54,718.59	7.46	0.05	0.97	54,728.39	7.88	0.06	
54,718.73	7.55	0.05	0.97	54,729.03	7.58	0.06	
54,719.22	8.26	0.08	0.97	54,729.24	7.79	0.06	
54,719.29	7.81	0.06	0.97	54,731.07	7.20	0.05	
54,719.36	8.05	0.07	0.97	54,731.36	7.58	0.06	
54,719.43	7.86	0.06	0.97	54,733.61	8.03	0.07	
54,719.50	8.05	0.07	0.97	54,735.09	7.59	0.06	
54,719.58	7.72	0.06	0.98	54,737.21	7.26	0.05	
54,719.65	7.94	0.06	0.97	54,737.63	6.93	0.04	
54,720.07	8.36	0.08	0.97	54,738.06	6.94	0.04	
54,720.14	8.04	0.07	0.97	54,738.13	7.00	0.04	
54,720.21	7.84	0.06	0.97	54,738.20	7.09	0.04	
54,720.28	8.16	0.07	0.97				

54,761.27

54,761.34

54,761.41

54,761.48

54,761.55

54,761.62

54,761.76

54,761.83

54,762.04

8.50

8.13

9.00

8.81

9.64

9.11

9.11

11.87

11.77

0.08

0.07

0.11

0.10

0.14

0.11

0.11 0.39

0.37

Table 10 SMEI Light-curve Data for Nova QYMus Continued

SMEI Light-curve Data for Nova QYMus Continued				(Continued)				
Date-Obs MJD	SMEI Mag	SMEI Error	r	Date-Obs MJD	SMEI Mag	SMEI Error	r	
54,738.34	6.95	0.04	0.97	54,762.18	9.29	0.12	0.97	
54,738.55	7.16	0.05	0.97	54,762.25	9.08	0.11	0.97	
54,739.54	7.13	0.04	0.97	54,762.32	11.02	0.26	0.97	
54,739.68	7.29	0.05	0.97	54,762.54	8.60	0.09	0.97	
54,743.63	8.24	0.07	0.97	54,762,75	9.76	0.15	0.96	
54,745.11	8.52	0.08	0.97	54,762.82	9.31	0.12	0.97	
54,745.39	7.96	0.07	0.95	54,763.03	9.77	0.15	0.97	
54,745.54	8.18	0.07	0.97	54,763,59	9.16	0.11	0.97	
54,746.38	8.50	0.08	0.97	54,763.66	9.68	0.14	0.97	
54,747.23	8.65	0.09	0.97	54,763,73	9.20	0.12	0.97	
54,748.14	8.28	0.08	0.97	54,764.09	7.64	0.06	0.97	
54,748.50	9.03	0.11	0.97	54,764,16	8.36	0.08	0.97	
54,748.64	10.54	0.21	0.97	54.764.30	8.60	0.09	0.97	
54,748.78	8.27	0.08	0.97	54 764 65	10.05	0.17	0.97	
54,749.48	8.19	0.07	0.97	54,771,70	9 59	0.14	0.96	
54,749.70	9.04	0.11	0.96	54 772 06	8 38	0.08	0.96	
54,750.05	8.02	0.07	0.97	54 772 41	8.43	0.08	0.97	
54,750.26	9.45	0.13	0.97	54 772 48	11 49	0.00	0.97	
54,750.68	8.20	0.07	0.96	54 772 62	9.13	0.11	0.97	
54,751.04	8.25	0.07	0.97	54 772 76	9.08	0.11	0.96	
54,751.11	9.76	0.15	0.97	54 772 83	8.36	0.08	0.96	
54,751.25	8.20	0.07	0.97	54 772 90	10.83	0.00	0.90	
54,751.32	7.69	0.06	0.97	54,773.05	8.04	0.24	0.94	
54,751.39	10.00	0.17	0.97	54,773.12	8.04	0.07	0.90	
54,751,74	8.24	0.07	0.97	54,773.10	9.00	0.11	0.90	
54,751,81	7.95	0.07	0.96	54,772.26	0.59	0.07	0.97	
54.752.02	7.72	0.06	0.97	54,773.20	9.50	0.14	0.97	
54.752.30	9.00	0.11	0.97	54,773.33	8.00 8.10	0.09	0.90	
54,752,66	8 65	0.09	0.96	54,773.40	0.10 0.10	0.07	0.90	
54.752.73	9.57	0.14	0.97	54,772,54	0.10	0.07	0.90	
54,753,01	8.51	0.08	0.97	54,772,61	7.80	0.00	0.90	
54,753,08	9.11	0.11	0.97	54,772.69	8.05 0.28	0.07	0.90	
54,753,15	8.24	0.07	0.97	54,775.08	9.28	0.12	0.90	
54 753 22	8.03	0.07	0.97	54,772.82	8.22	0.07	0.90	
54 753 29	8 59	0.09	0.97	54,773.82	10.59	0.22	0.90	
54 753 36	9.41	0.13	0.97	54,773.89	8.35	0.08	0.92	
54 753 58	8 68	0.09	0.97	54,774.05	8.25	0.07	0.96	
54 754 56	8.84	0.10	0.97	54,774.10	8.54	0.09	0.96	
54 754 63	8 59	0.09	0.97	54,774.18	8.00	0.07	0.96	
54 754 85	7.48	0.05	0.97	54,774.25	9.53	0.13	0.96	
54,755,13	9.01	0.05	0.92	54,774.32	7.91	0.06	0.96	
54 755 34	9.61	0.14	0.97	54,774.39	8.31	0.08	0.96	
54 755 48	11.07	0.14	0.97	54,774.46	8.35	0.08	0.97	
54 755 60	8.03	0.41	0.96	54,774.53	9.67	0.14	0.97	
54,755.09	8 78	0.10	0.96	54,774.88	11.53	0.33	0.95	
54 756 40	10 11	0.17	0.96	54,774.95	9.23	0.12	0.92	
54 756 54	0.11	0.11	0.96					
54 756 61	874	0.00	0.90					
54,750.01 54 756 68	0./4	0.09	0.90					
54,750.00 54,756,75	9 60	0.40	0.90		Table 11			
54,130.13 54,757.02	0.00	0.09	0.97	SME	EI Light-curve Data for	Nova V5583 Sgr		
54,151.05 54 757 17	0.00	0.09	0.97		6			
34,/3/.1/ 54,757,22	8.33	0.08	0.97	Date-Obs MJD	SMEI Mag	SMEI Error	r	
54,131.32 54,757,20	9.82	0.15	0.90	55 044 16	8.60	0.10	0.23	
54,131.39	8.90	0.10	0.96	55 044 23	8 29	0.08	0.23	
34,/00.84	11.27	0.30	0.95	55,044.25	0.27	0.00	0.19	

0.95	55,044.23	8.29	0.08	0.19	
0.97	55,044.37	8.21	0.08	0.21	
0.97	55,044.44	8.65	0.10	0.23	
0.97	55,044.58	9.88	0.17	0.20	
0.97	55,048.67	8.01	0.07	0.24	
0.97	55,048.74	7.83	0.07	0.24	
0.97	55,048.81	7.77	0.07	0.21	
0.97	55,048.88	7.64	0.06	0.23	
0.96	55,048.96	7.60	0.06	0.26	
0.97	55,049.10	7.69	0.06	0.26	

Table 11 (Continued)

SMEI Error 0.06 0.05 0.04 0.03 0.02 0.02 0.02 0.01 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02 0.02

 Table 12

 SMEI Light-curve Data for Nova RS Oph

5.30

0.02

Dete Ohe MID	SMEL M	CMEL E		Date-Obs MID	SMEI Mag
	SWIEL Mag	SMEI Error	<u>r</u>	53 775 77	7 70
55,049.24	7.48	0.06	0.29	53,775,84	7.79
55,049.31	7.67	0.06	0.33	53 776 26	7.42
55,049.38	7.32	0.05	0.31	53 778 03	6.18
55,049.45	7.28	0.05	0.31	53 778 10	5.66
55,049.66	7.20	0.05	0.36	53 778 17	5 30
55,049.73	7.18	0.05	0.37	53 778 24	4.93
55,049.87	7.17	0.05	0.38	53 778 31	4.50
55,049.94	7.04	0.05	0.39	53 778 38	4.59
55,050.01	7.08	0.05	0.38	52 779 45	4.44
55,050.08	6.94	0.05	0.44	52,778.43	4.40
55,050.16	7.03	0.05	0.40	52,778,50	4.50
55,050.23	6.98	0.05	0.42	52,778.66	4.20
55,050.44	7.13	0.05	0.39	55,778.00	4.09
55,050.72	7.28	0.05	0.34	53,778.73	3.97
55,050.79	7.34	0.05	0.32	53,778.80	3.93
55,051.00	7.37	0.06	0.31	53,//8.8/	3.93
55,051.14	7.57	0.06	0.27	53,778.94	3.87
55,051.21	7.55	0.06	0.28	53,779.01	3.93
55,051.29	7.65	0.06	0.24	53,779.09	3.98
55,051.36	7.58	0.06	0.27	53,779.16	4.03
55,051.43	7.57	0.06	0.28	53,779.23	4.12
55,051.50	7.54	0.06	0.27	53,779.30	4.16
55.051.57	7.62	0.06	0.26	53,779.37	4.23
55,051,64	7.58	0.06	0.26	53,779.44	4.26
55 051 71	7.61	0.06	0.26	53,779.51	4.29
55,051.78	7.51	0.06	0.20	53,779.58	4.39
55 051 85	7.55	0.00	0.24	53,779.65	4.38
55,051,02	7.50	0.00	0.25	53,779.72	4.46
55,051.92	7.05	0.00	0.25	53,779,79	4.47
55,052,06	7.55	0.00	0.23	53,779.86	4.51
55,052.00	7.50	0.00	0.24	53,779.93	4.56
55,052.15	7.02	0.00	0.24	53 780 00	4 59
55,052.20	7.02	0.00	0.25	53 780 07	4 59
55,052.27	7.44	0.06	0.26	53 780 14	4 59
55,052.34	7.55	0.06	0.25	53 780 21	4 63
55,052.41	/.56	0.06	0.25	53 780 28	4 63
55,052.48	7.56	0.06	0.27	53,780,26	4.62
55,052.55	7.48	0.06	0.28	53 780 43	4.62
55,052.70	7.50	0.06	0.29	53,780.45	4.67
55,052.77	7.50	0.06	0.30	53,780.50	4.07
55,053.26	7.91	0.07	0.29	53,780.57	4.00
55,053.33	7.84	0.07	0.28	52,780.04	4.75
55,053.40	7.93	0.07	0.27	53,780.71	4.72
55,053.47	7.98	0.07	0.25	<i>JJ</i> ,/00./0 52 780 85	4./3
55,053.54	7.91	0.07	0.25	<i>52,700.02</i>	4./1
55,053.61	7.91	0.07	0.24	55,780.92	4./3
55,053.68	8.00	0.07	0.21	55,/81.27	4.91
55,053.75	8.11	0.08	0.20	55,/81.54	4.88
55,053.82	7.90	0.07	0.19	55,781.41	4.85
55,053.89	8.12	0.08	0.17	53,781.48	4.90
55,053.96	8.29	0.08	0.14	53,781.55	4.96
55,054.04	8.09	0.08	0.16	53,781.63	4.99
55,054.11	7.92	0.07	0.18	53,781.70	4.95
55,054.18	8.34	0.09	0.15	53,781.77	5.00
55,054.25	8.19	0.08	0.16	53,781.84	5.00
55,054.32	8.05	0.08	0.19	53,781.91	4.93
55,054.39	7.88	0.07	0.22	53,781.98	5.02
55,054.53	7.85	0.07	0.24	53,782.05	5.00
55,054.60	8.34	0.09	0.24	53,782.12	5.11
55,054.67	8.21	0.08	0.26	53,782.19	5.06
				53,782.26	5.13
				53,782.75	5.19
				53,782.82	5.26
				53,782.90	5.25

53,782.97

Date-Obs MJD

53,783.04

53,783.11 53,783.18

53,783.25

53,783.32

53,783.39

53,783.53

53,783.60

53,783.67

53,783.74

53,783.81

53,783.88 53,783.95

53,784.02

53,784.10

53,784.17 53,784.24

53,784.31

53,784.38

53,784.45

53,784.52

53,784.59

53,784.66

53,784.73

53,784.80

53,784.87

53,784.94

53,785.01

53,785.08

53,785.15

53,785.22

53,785.29

53,785.37

53,785.44

53,785.51

53,785.58

53,785.65

53,785.72

53,785.79

53,785.86

53,785.93

53,786.00

53,786.07

53,786.14

53,786.21

53,786.28

53,786.35

53,786.42

53,786.49

53,786.56

53,786.78

53,786.85

53,786.92

53,786.99

53,787.06

53,787.20

53,787.27

53,787.34

53,787.41

53,787.48

53,787.55

53,787.69

53,787.76

53,787.83

Table 12 (Continued)

HOUNSELL	ΕT	AL.
----------	----	-----

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.01

0.02

0.02

0.01

Table 12 Continued)		Table 12 (Continued)		
SMEI Mag	SMEI Error	Date-Obs MJD	SMEI Mag	SMEI Error
5.29	0.02	53,787.91	6.08	0.03
5.30	0.02	53,787.98	6.00	0.03
5.33	0.02	53,790.16	6.24	0.03
5.32	0.02	53,790.37	6.11	0.03
5.36	0.02	53,790.45	6.21	0.03
5.27	0.02	53,794.47	6.23	0.03
5.34	0.02	53,794.68	6.05	0.03
5.33	0.02	53,794.75	6.10	0.03
5.35	0.02	53,794.82	6.46	0.03
5.38	0.02	53,794.89	6.48	0.03
5.36	0.02	53,794.96	6.15	0.03
5.38	0.02			
5.38	0.02			
5.37	0.02			
5.47	0.02			
5.47	0.02			
5.38	0.02		Table 13	
5.56	0.02	SMEI L	ight-curve Data for Nova V1280 Sco	
5.45	0.02	Date-Obs MID	SMEI Mag	SMEL Error
5.46	0.02		Shirisi hing	
5.48	0.02	54,144.61	6.19	0.03
5.51	0.02	54,144.68	6.18	0.03
5.53	0.02	54,144.75	6.20	0.03
5.53	0.02	54,144.82	6.19	0.03
5.59	0.02	54,144.89	6.16	0.03
5.60	0.02	54,144.96	6.16	0.03
5.67	0.02	54,145.24	6.14	0.03
5.60	0.02	54,145.31	6.20	0.03
5.71	0.02	54,145.38	6.10	0.03
5.65	0.02	54,145.52	6.02	0.03
5.72	0.02	54,145.59	5.89	0.02
5.72	0.02	54,145.66	5.81	0.02
5.66	0.02	54,145.73	5.70	0.02
5.66	0.02	54,145.80	5.59	0.02
5.67	0.02	54,145.88	5.48	0.02
5.72	0.02	54,145.95	5.37	0.02
5.72	0.02	54,146.02	5.30	0.02
5.63	0.02	54,146.09	5.27	0.02
5.64	0.02	54,146.23	5.23	0.02
5.59	0.02	54,146.30	5.24	0.02
5.63	0.02	54,146.37	5.23	0.02
5.73	0.02	54,146.44	5.22	0.02
5.75	0.02	54,146.51	5.24	0.02
5.75	0.02	54,146.58	5.22	0.02
5.77	0.02	54,146.65	5.23	0.02
5.77	0.02	54,146.72	5.14	0.02
5.75	0.02	54,146.79	4.75	0.01
5.74	0.02	54,146.86	4.37	0.01
5 70	0.02	54,146.93	4.12	0.01

5.67	0.02	54,145.00	5.40
5.72	0.02	54,145.95	5.37
5.72	0.02	54,146.02	5.30
5.63	0.02	54,146.09	5.27
5.65	0.02	54,146.23	5.23
5 50	0.02	54,146.30	5.24
5.63	0.02	54,146.37	5.23
5.05	0.02	54,146.44	5.22
5.75	0.02	54,146.51	5.24
5.75	0.02	54,146.58	5.22
5.75	0.02	54,146.65	5.23
5.77	0.02	54,146.72	5.14
5.77	0.02	54,146.79	4.75
5.75	0.02	54,146.86	4.37
5.74	0.02	54,146.93	4.12
5.79	0.02	54,147.00	4.05
5.78	0.02	54,147.07	4.00
5.77	0.02	54,147,15	4.00
5.83	0.02	54.147.22	4.02
5.80	0.02	54,147,29	4.04
5.84	0.02	54.147.36	4.14
5.80	0.02	54,147,43	4 13
5.82	0.02	54.147.50	4.19
5.88	0.03	54 147 57	4.25
5.93	0.03	54 147 64	4 41
5.91	0.03	54 147 71	4 54
5.98	0.03	54 147 78	4 68
5.89	0.03	54 147 85	4.00
6.00	0.03	54 147 02	4.85
6.09	0.03	54 148 06	4.85
5.95	0.03	54 148 13	4.73
		54,140.15	4.73

Table 13

SMEI Error 0.02 0.02 0.01 0.01 0.01 0.01 0.02

(Continued)			(Continued)		
Date-Obs MJD	SMEI Mag	SMEI Error	Date-Obs MJD	SMEI Mag	
54,148.27	4.35	0.01	54,153.07	4.90	
54,148.34	4.23	0.01	54,153.14	4.85	
54,148.41	4.24	0.01	54,153.21	4.77	
54,148.49	4.25	0.01	54,153.28	4.77	
54,148.56	4.29	0.01	54,153.35	4.77	
54,148.63	4.33	0.01	54,153.42	4.80	
54,148.70	4.41	0.01	54,153.49	4.83	
54,148.77	4.49	0.01	54,153.56	4.84	
54,148.84	4.56	0.01	54,153.64	4.86	
54,148,91	4.64	0.01	54.153.71	4.85	
54.148.98	4.68	0.01	54,153.78	4.86	
54,149.05	4.71	0.01	54,153,85	4.88	
54,149,12	4.71	0.01	54,153.92	4.87	
54,149,19	4 74	0.01	54 154 06	4 84	
54,149,26	4 76	0.01	54 154 13	4 84	
54,149,33	4 79	0.01	54,154,20	4 83	
54 149 40	4 81	0.01	54 154 27	4.83	
54,149,47	4 80	0.01	54,154,34	4 87	
54 149 54	4 77	0.01	54 154 41	4 88	
54 149 61	4.77	0.01	54,154,48	4.86	
54 149 68	4.74	0.01	54,154,40	4.00	
54 149 76	4.70	0.01	54 154 62	4.86	
54,149.83	4.56	0.01	54 154 69	4.50	
54,149.00	4.50	0.01	54 154 76	4.80	
54,149.90	4.40	0.01	54,154.70	4.09	
54,149.97	4.25	0.01	54 154 01	4.95	
54,150.04	4.25	0.01	54 154 08	4.98	
54,150.11	4.10	0.01	54,154.98	4.97	
54,150.16	4.13	0.01	54,155.05	4.98	
54,150.25	4.15	0.01	54 155 10	4.08	
54,150.52	4.15	0.01	54 155 26	4.98	
54,150.59	4.22	0.01	54,155,20	4.90	
54,150.55	4.44	0.01	54 155 40	4.97	
54,150.00	4.37	0.01	54 155 47	4.98	
54,150.07	4.74	0.01	54 155 54	5.04	
54,150.74	4.92	0.02	54,155,54	5.04	
54,150.81	4.94	0.02	54,155.69	5.00	
54,150.00	4.89	0.02	54,155,08	5.10	
54,150.95	4.03	0.02	54 155 92	5.00	
54,151.05	4.01	0.01	54 155 80	5.05	
54,151.10	4.81	0.01	54,155,06	5.05	
54,151.17	4.80	0.02	54,156,02	5.02	
54,151.24	4.85	0.02	54,150.05	5.03	
54,151.51	4.80	0.02	54,156,19	5.03	
54,151.50	4.80	0.01	54,156,25	5.02	
54,151.45	4.75	0.01	54,156,22	5.03	
54,151.52	4.08	0.01	54,156,52	5.04	
54,151.59	4.07	0.01	54,156,91	5.04	
54,151.00	4.08	0.01	54,150.81	5.03	
54,151.75	4.71	0.01	54,157,22	5.04	
54,151.80	4.71	0.01	54,157,25	5.06	
54,151.87	4.71	0.01	54,157.30	5.11	
54,151.94	4.69	0.01	54,157.37	5.14	
54,152.22	4.78	0.01	54,157.45	5.15	
54,152.30	4.75	0.01	54,157.52	5.15	
54,152.37	4.77	0.01	54,157.59	5.16	
54,152.44	4.77	0.01	54,157.66	5.20	
54,152.51	4.73	0.01	54,157.73	5.22	
54,152.58	4.65	0.01	54,157.80	5.25	
54,152.65	4.70	0.01	54,157.87	5.25	
54,152.72	4.75	0.01	54,157.94	5.29	
54,152.79	4.80	0.01	54,158.01	5.30	
54,152.86	4.87	0.02	54,158.08	5.32	
54,152.93	4.89	0.02	54,158.15	5.35	
54,153.00	4.92	0.02	54,158.22	5.40	

Table 13
(Continued)

	(Continued)	
Date-Obs MJD	SMEI Mag	SMEI Error
54,158.29	5.42	0.02
54,158.36	5.44	0.02
54,158.43	5.48	0.02
54,158.50	5.49	0.02
54,158.57	5.53	0.02
54,158.64	5.51	0.02
54,158.72	5.53	0.02
54,158.79	5.57	0.02
54,158.86	5.55	0.02
54,158.93	5.60	0.02
54,159.28	5.63	0.02
54,159.35	5.73	0.02
54,159.42	5.81	0.02
54,159.49	5.87	0.02
54,159.56	5.92	0.02
54,159.63	5.95	0.03
54,159.70	5.94	0.03
54,159.77	5.95	0.03
54,159.84	5.99	0.03
54,159.98	6.00	0.03
54,160.13	6.07	0.03
54,160.27	6.19	0.03
54,160.41	6.17	0.03
54,160.48	6.18	0.03
54,160.55	6.22	0.03
54,160.62	6.23	0.03
54,160.69	6.24	0.03
54,160.76	6.26	0.03

HOUNSELL ET AL

Table 14 (Continu

Date-Obs MJD	SMEI Mag	SMEI Error
54,162.95	7.57	0.05
54,163.02	7.80	0.06
54,163.09	7.60	0.05
54,163.16	7.69	0.06
54,163.23	7.79	0.06
54,163.30	7.61	0.05
54,163.37	8.21	0.07
54,163.44	8.19	0.07
54,163.51	8.04	0.06
54,163.58	8.29	0.07
54,163.65	8.10	0.07
54,163.72	8.22	0.07
54,163.79	8.31	0.07
54,163.87	8.69	0.09
54,163.94	9.08	0.10
54,164.01	8.86	0.09
54,164.08	9.78	0.14
54,164.15	9.41	0.12
54,164.22	10.02	0.16
54,164.29	9.72	0.14
54,164.36	11.18	0.27
54,164.43	10.21	0.17
54,164.71	10.70	0.22

Table 15 SMEI Light-curve Data for Nova V598 Pup

		Date-Obs MJD	SMEI Mag	SMEI Error
bla 14		54,248.82	9.60	0.14
010 14 · Nova V1280 Sco	Continued	54,254.47	7.54	0.05
11010 11200 500	Continued	54,254.61	6.36	0.03
IEI Mag	SMEI Error	54,254.68	6.06	0.03
6.25	0.03	54,254.75	5.88	0.03
6.27	0.03	54,254.82	5.57	0.02
6 34	0.03	54,254.96	4.94	0.02
6 3 9	0.03	54,255.03	5.08	0.02
6.42	0.03	54,255.10	5.24	0.02
6 35	0.03	54,255.17	5.13	0.02
6.45	0.03	54,255.24	5.05	0.02
6 4 9	0.03	54,255.31	4.77	0.02
6.53	0.03	54,255.38	4.72	0.01
6.63	0.03	54,255.45	4.62	0.01
6.66	0.03	54,255.52	4.49	0.01
6.61	0.03	54,255.59	4.34	0.01
6.69	0.03	54,255.67	4.38	0.01
6.70	0.04	54,255.74	4.29	0.01
6.05	0.04	54,255.81	4.04	0.01
6.06	0.04	54,255.88	4.06	0.01
7.03	0.04	54,255.95	3.90	0.01
7.05	0.04	54,256.02	3.97	0.01
7.10	0.04	54,256.09	3.92	0.01
7.12	0.04	54,256.16	3.87	0.01
7.20	0.04	54,256.23	3.68	0.01
7.14	0.04	54,256.30	3.62	0.01
7.20	0.04	54,256.65	3.49	0.01
7.30	0.05	54,256.72	3.50	0.01
7.32	0.05	54,256.79	3.67	0.01
7.55	0.05	54,256.94	3.48	0.01
1.33 7.55	0.05	54,257.29	3.46	0.01
1.55	0.05	54,257.36	3.57	0.01
1.55	0.05	54,257.43	3.48	0.01

Table 1 SMEI Light-curve Data for Nor

SMEI

Date-Obs MJD 54,160.83

54,160.90

54,160.97

54,161.04

54,161.11

54,161.18

54,161.40

54,161.47

54,161.54

54,161.61

54,161.68

54,161.75

54,161.82

54,161.89

54,161.96

54,162.03 54,162.10

54,162.17

54,162.24

54,162.31

54,162.38

54,162.45

54,162.52

54,162.60 54,162.67

54,162.74

54,162.81 54,162.88

Table 15 (Continued)

Date-Obs MJD	SMEI Mag	SMEI Error
54,257.50	3.52	0.01
54,257.57	3.53	0.01
54,257.64	3.57	0.01
54,257.71	3.62	0.01
54,257.78	3.61	0.01
54,257.99	3.81	0.01
54,258.06	3.82	0.01
54,258.13	3.84	0.01
54,258.20	5.84 2.80	0.01
54,258,25	3.09	0.01
54 258 42	3.99	0.01
54 258 49	4 04	0.01
54.258.56	4.19	0.01
54,258.63	4.08	0.01
54,258.70	4.10	0.01
54,258.77	4.15	0.01
54,258.98	4.27	0.01
54,259.05	4.28	0.01
54,259.12	4.33	0.01
54,259.19	4.34	0.01
54,259.26	4.41	0.01
54,259.33	4.43	0.01
54,279.71	7.04	0.04
54,279.79	7.07	0.04
54,279.86	6.87	0.04
54,280.07	7.10	0.04
54,280,21	7.13	0.03
54 280 28	6 98	0.04
54.280.35	7.19	0.05
54.280.42	7.18	0.05
54,280.49	7.32	0.05
54,280.56	7.19	0.05
54,280.63	7.27	0.05
54,280.70	7.31	0.05
54,280.77	7.39	0.05
54,280.84	7.14	0.04
54,281.34	7.23	0.05
54,281.41	7.81	0.06
54,281.48	7.91	0.06
54,281.55	8.20	0.07
54,281.02	7.84	0.06
54,281.09	7.54	0.03
54 282 04	9.53	0.00
54 282 18	8 63	0.09
54.282.32	10.65	0.22
54.282.54	9.38	0.12
54,282.68	9.57	0.14
54,282.82	9.28	0.12
54,282.89	11.51	0.33
54,283.38	7.16	0.05
54,283.45	7.19	0.05
54,283.52	8.90	0.10
54,283.59	10.36	0.19
54,283.66	7.03	0.04
54,283.73	7.18	0.05
54,285.80	6.85	0.04
54,280.84	/.40	0.05
54,200.91	1.33 7 75	0.05
54 287 12	8.43	0.00
54,287,19	7 96	0.08
,=0//12/	1.20	5.07

HOUNSELL	EТ	AL.
1001010111		

Table 15 (Continued)

(continued)		
Date-Obs MJD	SMEI Mag	SMEI Error
54,287.26	8.24	0.07
54,287.33	8.02	0.07
54,287.40	8.23	0.07
54,287.47	8.27	0.08
54,287.54	8.02	0.07
54,287.61	9.09	0.11
54,287.68	8.16	0.07
54,287.75	8.21	0.07
54,288.18	7.97	0.07
54,288.25	7.95	0.06
54,288.32	7.64	0.06
54,288.46	7.50	0.05
54,288.53	8.01	0.07
54,288.60	8.94	0.10
54,288.67	8.71	0.09
54,288.74	7.76	0.06
54,288.81	7.72	0.06
54,288.88	7.64	0.06
54,288.95	7.92	0.06
54,289.16	8.46	0.08
54,289.23	8.36	0.08
54,289.31	8.52	0.08
54,289.38	7.90	0.06

 Table 16

 SMEI Light-curve Data for Nova V598 Pup Continued

Date-Obs MJD	SMEI Mag	SMEI Error
54,289.45	8.44	0.08
54,289.52	8.73	0.09
54,289.66	8.12	0.07
54,289.73	7.96	0.07
54,289.80	7.81	0.06
54,289.87	8.31	0.08
54,289.94	7.99	0.07
54,290.01	7.67	0.06
54,290.08	7.73	0.06
54,290.15	7.76	0.06
54,292.90	8.37	0.08
54,292.97	7.67	0.06
54,293.04	7.73	0.06
54,293.11	7.72	0.06
54,293.18	7.69	0.06
54,293.25	7.83	0.06
54,293.33	7.32	0.05
54,293.40	7.81	0.06
54,293.47	7.77	0.06
54,293.54	7.78	0.06
54,293.61	7.48	0.05
54,293.68	7.53	0.05
54,293.75	7.93	0.06
54,293.82	7.89	0.06
54,293.89	7.55	0.05
54,293.96	8.06	0.07
54,297.06	12.19	0.45
54,297.84	7.82	0.06
54,297.91	7.41	0.05
54,297.98	8.18	0.07
54,298.05	8.76	0.09
54,298.12	7.53	0.05
54,298.19	7.23	0.05

Table 16

SMEI Error 0.05 0.10 0.28 1.38 0.08 0.03 0.10 0.05 0.07 0.03 0.03 0.05 0.03 0.03 0.03 0.03 0.03 0.03 0.03 0.03 0.11 0.03 0.03 0.04 0.05 0.06 0.08 0.05 0.08 0.06 0.09 0.04 0.26 0.08 0.08 0.05 0.05 0.06 0.08 0.06 0.10 0.09 0.05 0.06 0.14 1.71 0.08 0.09 0.25 0.11 0.09 0.14 0.08 0.11 0.07 0.06 0.08 0.10 0.08 0.08 0.09 0.08 0.09 0.08

(Continued)		(Continued)		
Date-Obs MJD	SMEI Mag	SMEI Error	Date-Obs MJD	SMEI Mag
54,298.26	12.24	0.46	54,311.45	7.49
54,298.33	7.75	0.06	54,311.52	8.81
54,298.40	7.92	0.06	54,311.59	11.19
54,298.47	7.97	0.07	54,311.74	14.67
54,298.54	7.92	0.06	54,311.81	8.37
54,298.62	6.70	0.04	54,311.88	5.97
54,298.76	8.19	0.07	54,312.02	8.84
54,298.83	7.06	0.04	54,312.09	7.17
54,303.20	13.64	0.86	54,312.16	8.24
54,306.73	7.77	0.06	54,312.23	6.06
54,306.80	9.09	0.11	54,312.30	6.15
54,306.87	7.95	0.06	54,312.37	7.36
54,306.94	8.17	0.07	54,312.44	6.40
54,307.01	8.71	0.09	54,312.51	5.95
54,307.15	9.80	0.15	54,312.58	5.97
54,307.22	9.41	0.13	54,312.65	5.91
54,307.29	9.06	0.11	54,312.72	5.90
54,307.36	8.29	0.08	54,312.79	6.40
54,307.43	9.76	0.15	54,312.86	6.40
54,307.50	8.76	0.09	54,312.93	6.02
54,307.57	9.11	0.11	54,313.29	9.20
54,307.64	9.55	0.13	54,313.36	6.45
54,307.71	10.08	0.17	54,313.43	6.29
54,307.78	9.05	0.11	54,313.50	7.06
54,308.00	10.45	0.20	54,313.57	7.43
54,308.07	8.37	0.08	54,313.78	7.84
54,308.14	7.68	0.06	54,313.85	8.40
54,308.28	9.97	0.16	54,313.92	7.28
54,308.35	7.61	0.06	54,313.99	8.49
54,308.42	7.39	0.05	54,314.06	7.90
54,308.49	10.21	0.18	54,314.13	8.64
54,308.56	7.47	0.05	54,314.20	7.08
54,308.63	9.67	0.14	54,314.27	11.04
54,308.70	6.45	0.03	54,314.35	8.41
54,308.84	8.93	0.10	54,314.42	8.35
54,308.92	7.53	0.05	54,314.49	7.53
54,308.98	7.94	0.06	54,314.63	7.54
54,309.27	8.23	0.07	54,314.70	7.92
54,309.34	8.28	0.08	54,314.77	8.38
54,309.41	8.00	0.07	54,314.84	7.59
54,309.48	7.79	0.06	54,314.91	8.91
54,309.55	7.87	0.06	54,314.98	8.69
54,309.62	7.84	0.06	54,315.05	7.53
54,309.69	10.94	0.25	54,315.12	7.93
54,309.76	7.68	0.06	54,315.19	9.62
54,309.83	8.33	0.08	54,315.40	15.13
54,309.90	8.21	0.07	54,315.54	8.31
54,309.97	8.68	0.09	54,315.61	8.73
54,310.04	8.58	0.09	54,315.76	10.90
54,310.11	8.65	0.09	54,315.83	9.13
54,310.18	8.61	0.09	54,315.90	8.61
54,310.25	8.31	0.08	54,315.97	9.68
54,310.32	9.13	0.11	54,316.04	8.39
54,310.39	8.73	0.09	54,316.11	9.14
54,310.47	9.05	0.11	54,316.18	8.06
54,310.61	9.08	0.11	54,316.25	7.62
54,310.68	8.04	0.07	54,316.32	8.47
54,310.75	9.22	0.12	54,316.39	8.95
54,310.82	11.30	0.30	54,316.53	8.38
54,310.89	9.57	0.14	54,316.60	8.39
54,310.96	7.92	0.06	54,316.67	8.67
54,311.17	8.71	0.09	54,316.74	8.33
54,311.31	8.49	0.08	54,316.81	8.64
54,311.38	9.26	0.12	54,316.88	8.35

Table 16

	(Continued)			
Date-Obs MJD	SMEI Mag	SMEI Error	Date-Obs MJD	
54,317.10	9.54	0.13	54,321.89	
54,317.31	10.29	0.19	54,321.96	
54,317.38	10.89	0.25	54,322.10	
54,317.45	8.62	0.09	54,322.18	
54,317.52	7.82	0.06	54,322.25	
54,317.59	9.60	0.14	54,322.32	
54,317.66	8.23	0.07	54,322.39	
54,317.73	8.28	0.08	54,322.46	
54,317.80	7.75	0.06	54,322.53	
54,317.87	9.46	0.13	54,322.60	
54,317.94	8.10	0.07	54,322.67	
54,318.01	7.95	0.06	54,322.74	
54,318.08	8.17	0.07	54,322.95	
54,318.15	7.26	0.05	54,323.23	
54,318.22	11.12	0.27	54,323.37	
54,318.30	8.48	0.08	54,323.45	
54,318.44	8.41	0.08	54,323.59	
54,318.51	13.05	0.66	54,323.66	
54,318.58	9.05	0.11	54,323.73	
54,318.65	8.41	0.08	54,323.80	
54,318.72	8.69	0.09	54,323.87	
54,318.79	11.58	0.34	54,323.94	
54,318.86	8.96	0.10	54,324.01	
54,318.93	7.52	0.05	54,324.08	
54,319.00	8.52	0.08	54,324.15	
54,319.07	9.98	0.16	54,324.29	
54,319.21	8.34	0.08	54,324.57	
54,319.35	8.52	0.08	54,324.64	
54,319.42	7.46	0.05	54,324.72	
54,319.49	8.15	0.07	54,324.79	
54,319.57	11.33	0.30	54,324.86	
54,319.64	8.42	0.08	54,324.93	
54,319.71	7.40	0.05	54,325.14	
54,319.78	6.81	0.04	54,325.21	
54,319.85	8.53	0.08	54,325.28	
54,319.92	8.36	0.08	54,325.35	
54,319.99	8.22	0.07	54,325.42	
54,320.27	8.90	0.10	54,325.63	
54,320.34	8.23	0.07	54,325.70	
54,320.41	8.18	0.07	54,325.77	
54,320.48	7.79	0.06	54,325.84	
54,320.55	8.14	0.07	54,325.91	
54,320.62	8.51	0.08	54,325.98	
54,320.69	8.67	0.09	54,326.06	
54,320.76	8.32	0.08	54,326.20	
54,320.84	9.45	0.13	54,326.27	
54,320.91	8.72	0.09	54,326.34	
54,321.12	8.64	0.09	54,326.41	
54,321.19	7.80	0.06	54,326.48	
			54,326.55	
			54,326.62	

 Table 17

 SMEI Light-curve Data for Nova V598 Pup Continued

Date-Obs MJD	SMEI Mag	SMEI Error
54,321.26	8.26	0.07
54,321.47	8.56	0.09
54,321.54	8.73	0.09
54,321.61	9.40	0.13
54,321.68	9.08	0.11
54,321.75	9.88	0.16
54,321.82	9.48	0.13

Date-Obs MJD	SMEI Mag	SMEI Error
54,321.89	6.84	0.04
54,321.96	9.38	0.12
54,322.10	8.04	0.07
54,322.18	11.40	0.31
54,322.25	8.53	0.08
54,322.32	9.09	0.11
54,322.39	7.61	0.06
54,322.46	6.88	0.04
54,322.53	7.10	0.04
54,322.60	8.16	0.07
54,522.07	7.18	0.03
54,322.74	7.75	0.00
54 323 23	8.04	0.00
54.323.37	9.96	0.16
54,323,45	9.68	0.14
54,323.59	8.15	0.07
54,323.66	6.88	0.04
54,323.73	9.08	0.11
54,323.80	7.46	0.05
54,323.87	7.86	0.06
54,323.94	7.62	0.06
54,324.01	7.97	0.07
54,324.08	8.34	0.08
54,324.15	7.64	0.06
54,324.29	7.98	0.07
54,324.57	7.95	0.06
54,524.04 54,224.72	7.57	0.05
54,324.72	7.45	0.05
54 324 86	8.51	0.05
54,324,93	9.02	0.11
54.325.14	7.23	0.05
54,325.21	7.67	0.06
54,325.28	10.04	0.17
54,325.35	9.85	0.15
54,325.42	7.43	0.05
54,325.63	8.67	0.09
54,325.70	8.47	0.08
54,325.77	7.99	0.07
54,325.84	9.75	0.15
54,325.91	8.30	0.08
54,325.98	8.70	0.09
54,326.06	9.49	0.13
54,520.20 54,326.27	0.47 8 20	0.03
54,326.37	10.41	0.07
54 326 41	8 53	0.20
54,326,48	9 59	0.14
54,326.55	7.47	0.05
54,326.62	9.28	0.12
54,326.76	8.02	0.07
54,326.83	8.39	0.08
54,326.90	9.02	0.11
54,326.97	8.01	0.07
54,327.25	8.02	0.07
54,327.33	9.54	0.13
54,327.40	9.27	0.12
54,327.47	7.47	0.05
54,327.54	9.75	0.15
54,527.01	9.44	0.13
54,527.08	9.20 7 76	0.11
54 327 82	0 10	0.00
5 1,521.02	1.17	0.11

Table 17 (Continued)

Table 17

(Continued)		
Date-Obs MJD	SMEI Mag	SMEI Error
54,327.89	9.00	0.10
54,327.96	9.17	0.11
54,328.03	8.78	0.09
54,328.10	8.19	0.07
54,328.17	8.40	0.08
54,328.24	8.58	0.09
54,328.31	8.39	0.08
54,328.45	10.60	0.22
54,328.67	9.57	0.14
54,328.74	9.56	0.13
54,328.81	8.16	0.07
54,328.88	8.27	0.08
54,328.95	9.24	0.12
54,329.02	7.62	0.06
54,329.09	8.58	0.09
54,329.16	8.87	0.10
54,329.23	9.06	0.11
54,329.37	8.51	0.08
54,329.44	10.23	0.18
54,329.51	8.49	0.08
54,329.58	9.00	0.10
54,329.65	8.43	0.08
54,329.72	9.09	0.11
54,329.79	9.20	0.11
54,329.87	11.10	0.27
54,329.94	8.55	0.09
54,330.01	8.21	0.07
54,330.08	11.81	0.38
54,330.15	8.35	0.08
54,330.29	8.54	0.09
54,330.36	8.33	0.08
54,330.43	7.37	0.05
54,330.50	7.79	0.06
54,330.57	9.18	0.11
54,330.64	8.78	0.09
54,330.71	7.83	0.06
54,330.85	9.85	0.15
54,330.99	8.70	0.09

Table 18SMEI Light-curve Data for Nova KT Eri

Date-Obs MJD	SMEI Mag	SMEI Error
55,146.99	9.51	0.15
55,148.12	8.44	0.09
55,148.19	8.05	0.08
55,148.26	7.83	0.07
55,148.33	7.50	0.06
55,148.40	7.28	0.05
55,148.47	7.01	0.05
55,148.54	6.84	0.04
55,148.62	6.67	0.04
55,148.69	6.52	0.04
55,148.76	6.35	0.04
55,148.83	6.17	0.03
55,148.90	5.94	0.03
55,148.97	6.02	0.03
55,149.25	5.78	0.03
55,149.32	5.74	0.03
55,149.39	5.61	0.03
55,149.46	5.53	0.02

Table 18 (Continued)		
Date-Obs MJD	SMEI Mag	SMEI Error
55,149.53	5.47	0.02
55,149.60	5.45	0.02
55,149.67	5.42	0.02
55,149.74	5.44	0.02
55,149.81	5.47	0.02
55,149.89	5.49	0.02
55,149.90 55,150.02	5.59	0.02
55 150 10	5.62	0.03
55,150,17	5.80	0.03
55,150.24	5.85	0.03
55,150.45	5.90	0.03
55,150.52	5.88	0.03
55,150.59	5.94	0.03
55,150.66	5.95	0.03
55,150.73	5.98	0.03
55,150.80 55,151.08	6.01	0.03
55 151 16	6.19	0.03
55 151 23	6.19	0.03
55,151.30	6.23	0.03
55,151.37	6.31	0.03
55,151.44	6.34	0.03
55,151.51	6.33	0.03
55,151.58	6.34	0.03
55,151.65	6.35	0.04
55,151.72	6.36	0.04
55,151.79	6.37	0.04
55,151.80	6.38	0.04
55 152 00	6.32	0.03
55,152.00	6.44	0.04
55,152,14	6.47	0.04
55,152.21	6.44	0.04
55,152.28	6.52	0.04
55,152.35	6.50	0.04
55,152.43	6.53	0.04
55,152.50	6.56	0.04
55,152.57	6.60	0.04
55,152.04 55,152.71	6.60	0.04
55 152 78	6.58	0.04
55,152.85	6.52	0.04
55,152.92	6.65	0.04
55,153.27	6.69	0.04
55,153.34	6.65	0.04
55,153.41	6.72	0.04
55,153.48	6.75	0.04
55,153.55	6.72	0.04
55,153.63	6.82	0.04
55,153.70 55,152,77	6.88	0.04
55 153 84	6.93	0.03
55 153 01	6.95	0.05
55.153.98	7.01	0.05
55,154.05	7.04	0.05
55,154.12	7.03	0.05
55,154.19	7.07	0.05
55,154.26	7.11	0.05
55,154.33	7.14	0.05
55,154.40	7.20	0.05
55,154.75	7.20	0.05
55,154.82	7.21	0.05
33,134.90	1.27	0.05

Table 18 (Continued)

Date-Obs MJD	SMEI Mag	SMEI Error
55 154 97	7 29	0.05
55,155,04	7.32	0.05
55,155,11	7.30	0.05
55,155.18	7.35	0.06
55,155.25	7.28	0.05
55,155.32	7.26	0.05
55,155.39	7.22	0.05
55,155.53	7.22	0.05
55,155.60	7.23	0.05
55,155.67	7.25	0.05
55,155.74	7.30	0.05
55,155.81	7.36	0.06
55,155.88	7.36	0.06
55,155.95	7.35	0.06
55,156.02	7.35	0.06
55,156.09	7.37	0.06
55,156.17	7.34	0.05
55,156.24	7.47	0.06
55,156.31	7.52	0.06
55,156.38	7.50	0.06
55,156.52	7.55	0.06
55,156.59	7.51	0.06
55,156.66	7.59	0.06
55,156.73	7.60	0.06
55,156.80	7.66	0.06
55,156.87	7.62	0.06
55,150.94	7.03	0.06
55,157.01	7.59	0.06
55,157.06	7.59	0.00
55 157 22	7.56	0.00
55 157 37	7.50	0.06
55 157 44	7.52	0.06
55,157,51	7.62	0.06
55,157,58	7.65	0.06
55,157.65	7.66	0.06
55,157.72	7.56	0.06
55,157.79	7.62	0.06
55,157.86	7.68	0.06
55,158.14	7.77	0.07
55,158.21	7.84	0.07
55,158.28	7.78	0.07
55,158.35	7.76	0.07
55,158.42	7.83	0.07
55,158.56	7.84	0.07
55,158.64	7.78	0.07
55,158.71	7.96	0.07
55,158.78	7.83	0.07
55,158.85	7.92	0.07
55,158.92	7.95	0.07
55,158.99	8.01	0.07
55,159.06	8.03	0.08
55,159.13	8.03	0.08
55,159.20 55,150,27	8.04	0.08
55,159.27	8.U3 0 1 4	0.08
55,159,54 55,150,41	ð.14 9 10	0.08
55,159.41 55,150,62	0.19	0.08
55,159.02 55,150,60	0.20 8 25	0.08
55 150 76	0.2J 8 77	0.08
55 159 84	8 26	0.08
55 159 91	8 23	0.00
55 159 98	8.17	0.08
55 160 26	8 21	0.08
22,100.20	0.21	0.00

Table 18

Date-Obs MJD	SMEI Mag	SMEI Error
55,160.33	8.24	0.08
55,160.40	8.23	0.08
55,160.47	8.20	0.08
55,160.54	8.23	0.08
55,160.61	8.35	0.09
55,160.68	8.31	0.09
55,160.82	8.35	0.09
55,160.89	8.29	0.08
55,160.96	8.29	0.08
55,161.03	8.23	0.08
55,161.11	8.35	0.09
55,161.18	8.33	0.09
55,162.23	8.27	0.08

REFERENCES

- Bode, M. F. 2010, AN, 331, 160
- Bode, M. F., & Evans, A. (ed.) 2008, Classical Novae (2nd ed; Cambridge: Cambridge Univ. Press)
- Buffington, A., Band, D. L., Jackson, B. V., Hick, P. P., & Smith, A. C. 2006, ApJ, 637, 880
- Buffington, A., Morrill, J. S., Hick, P. P., et al. 2007, Proc. SPIE, 6689, 66890B
- Chandrasekhar, S. 1931, ApJ, 74, 81
- Clover, J., Jackson, B. V., Buffington, A., et al. 2011, BAAS, 43, 257.02
- Darnley, M. J., Bode, M. F., Harman, D. J., et al. 2013, arXiv:1303.2711
- Darnley, M. J., Bode, M. F., Kerins, E., et al. 2006, MNRAS, 369, 257
- Darnley, M. J., Henze, M., Steele, I. A., et al. 2015, A&A, 580, A45
- Darnley, M. J., Ribeiro, V. A. R. M., Bode, M. F., Hounsell, R. A., & Williams, R. P. 2012, ApJ, 746, 61
- Darnley, M. J., Ribeiro, V. A. R. M., Bode, M. F., & Munari, U. 2011, A&A, 530, A70
- Darnley, M. J., Williams, S. C., Bode, M. F., et al. 2014, A&A, 563, L9
- Eyles, C. J., Harrison, R. A., Davis, C. J., et al. 2009, SoPh, 254, 387
- Eyles, C. J., Simnett, G. M., Cooke, M. P., et al. 2003, SoPh, 217, 319
- Gaposchkin, C. H. P. 1957, The Galactic Novae (New York, Amsterdam: Interscience, North-Holland)
- Goss, K. J. F., Karoff, C., Chaplin, W. J., Elsworth, Y., & Stevens, I. R. 2011, MNRAS, 411, 162
- Hachisu, I., & Kato, M. 2004, ApJL, 612, L57
- Helton, L. A., Woodward, C. E., Vanlandingham, K., & Schwarz, G. J. 2008, CBET, 1379, 1
- Henden, A., & Munari, U. 2008, IBVS, 5822, 1
- Henze, M., Darnley, M. J., Kabashima, F., et al. 2015, A&A, 582, L8
- Hernanz, M. 2005, in ASP Conf. Ser. 330, The Astrophysics of Cataclysmic Variables and Related Objects, ed. J.-M. Hameury, & J.-P. Lasota (San Francisco, CA: ASP), 265
- Hick, P., Buffington, A., & Jackson, B. V. 2005, Proc. SPIE, 5901, 340
- Hick, P., Buffington, A., & Jackson, B. V. 2007, Proc. SPIE, 6689, 66890C
- Hillman, Y., Prialnik, D., Kovetz, A., Shara, M. M., & Neill, J. D. 2014, MNRAS, 437, 1962
- Holdsworth, D. L., Rushton, M. T., Bewsher, D., et al. 2014, MNRAS, 438, 3483
- Hounsell, R., Bode, M. F., Hick, P., et al. 2010a, ATel, 2558, 1
- Hounsell, R., Bode, M. F., Hick, P. P., et al. 2010b, ApJ, 724, 480 (Paper I). Hounsell, R., Darnley, M. J., Harman, D. J., et al. 2011, ATel, 3373, 1
- Hounsell, R. A., Bode, M. F., Darnley, M. J., et al. 2012, in IAU Symp. 285, New Horizons in Time-domain Astronomy, ed. E. Griffin, R. Hanisch, & R. Seaman (Cambridge: Cambridge Univ. Press), 91
- Jackson, B. V., Buffington, A., Hick, P. P., et al. 2004, SoPh, 225, 177
- Liller, W. 2008, IAUC, 9004, 1
- Liller, W., Tabur, V., Williams, P., et al. 2008, IAUC, 8990, 2
- Lynch, D. K., Russell, R. W., Rudy, R. J., & Woodward, C. E. 2009, BAAS, 41, 491.25
- Lynch, D. K., Woodward, C. E., Geballe, T. R., et al. 2006, ApJ, 638, 987
- McLaughlin, D. B. 1960, in Stellar Atmospheres, ed. J. L. Greenstein (Chicago: Univ. Chicago Press), 585

- Munari, U., Dalla Via, G., Valisa, P., Dallaporta, S., & Castellani, F. 2007, CBET, 897, 1
- Munari, U., Siviero, A., Dallaporta, S., et al. 2011, NewA, 16, 209
- Naik, S., Banerjee, D. P. K., & Ashok, N. M. 2009, MNRAS, 394, 1551
- Nakano, S., Beize, J., Jin, Z.-W., et al. 2008a, IAUC, 8934, 1
- Nakano, S., Kadota, K., Waagen, E., et al. 2007a, IAUC, 8861, 2
- Nakano, S., Tago, A., Itagaki, K., & Kadota, K. 2007b, CBET, 1181, 1
- Nakano, S., Tago, A., Itagaki, K., et al. 2008b, IAUC, 8907, 1
- Nakano, S., Tago, A., Nishiyama, K., & Sakamoto, T. 2007c, IAUC, 8821, 1
- Nishiyama, K., Kabashima, F., Pojmanski, G., et al. 2009, IAUC, 9061, 1
- Page, K. L., Osborne, J. P., Evans, P. A., et al. 2010, MNRAS, 401, 121
- Pereira, A. J. S., McGaha, J. E., Young, J., & Rhoades, H. 2007, IAUC, 8895, 1
- Poggiani, R. 2008, Ap&SS, 315, 79
- Poggiani, R. 2010, NewA, 15, 170
- Pojmanski, G., Pilecki, B., & Szczygiel, D. 2005, AcA, 55, 275
- Schaefer, B. E., Pagnotta, A., LaCluyze, A. P., et al. 2011, ApJ, 742, 113
- Schwarz, G. J., Ness, J.-U., Osborne, J. P., et al. 2011, ApJS, 197, 31
- Shafter, A. W. 1997, ApJ, 487, 226
- Shafter, A. W. 2002, in AIP Conf. Proc. 637, Classical Nova Explosions, ed. M. Hernanz, & J. José (Melville, NY: AIP), 462
- Shafter, A. W., Rau, A., Quimby, R. M., et al. 2009, ApJ, 690, 1148
- Spreckley, S., & Stevens, I. R. 2007, in ASP Conf. Ser. 366, Transiting Extrapolar Planets Workshop, ed. C. Afonso, D. Weldrake, & Th. Henning (San Francisco, CA: ASP), 39

- Surina, F., Hounsell, R. A., Bode, M. F., et al. 2014, ApJ, 47, 107
- Swierczynski, E., Mikolajewski, M., Tomov, T., et al. 2010, in ASP Conf. Ser. 435, Binaries—Key to Comprehension of the Universe, ed. A. Prša & M. Zejda (San Francisco, CA: ASP), 297
- Tanaka, J., Nogami, D., Fujii, M., et al. 2011, PASJ, 63, 911
- Tarrant, N. J., Chaplin, W. J., Elsworth, Y., Spreckley, S. A., & Stevens, I. R. 2007, MNRAS, 382, L48
- Tarrant, N. J., Chaplin, W. J., Elsworth, Y., Spreckley, S. A., & Stevens, I. R. 2008a, A&A, 483, L43
- Tarrant, N. J., Chaplin, W. J., Elsworth, Y. P., Spreckley, S. A., & Stevens, I. R. 2008b, CoAst, 157, 92
- Tarrant, N. J., Chaplin, W. J., Elsworth, Y. P., Spreckley, S. A., & Stevens, I. R. 2008c, A&A, 492, 167
- Tomov, T., Mikolajewski, M., Brozek, T., et al. 2008, ATel, 1485, 1
- Walter, F. M., Battisti, A., Towers, S. E., Bond, H. E., & Stringfellow, G. S. 2012, PASP, 124, 1057
- Warner, B. 2008, in Properties of Novae: an Overview; Classical Novae, ed. M. F. Bode, & A. Evans (2nd ed.; Cambridge: Cambridge Univ. Press), 16 Warner, B., & Woudt, P. A. 2009, MNRAS, 397, 979
- Wesson, R., Barlow, M. J., Corradi, R. L. M., et al. 2008, ApJL, 688, L21
- Williams, R. 2012, AJ, 144, 98
- Wolf, W. M., Bildsten, L., Brooks, J., & Paxton, B. 2013, ApJ, 777, 136 Yamaoka, H. 2004, IAUC, 8380, 1
- Yamaoka, H., Ayani, K., Murakami, N., et al. 2007, CBET, 1181, 2
- Yaron, O., Prialnik, D., Shara, M. M., & Kovetz, A. 2005, ApJ, 623, 398